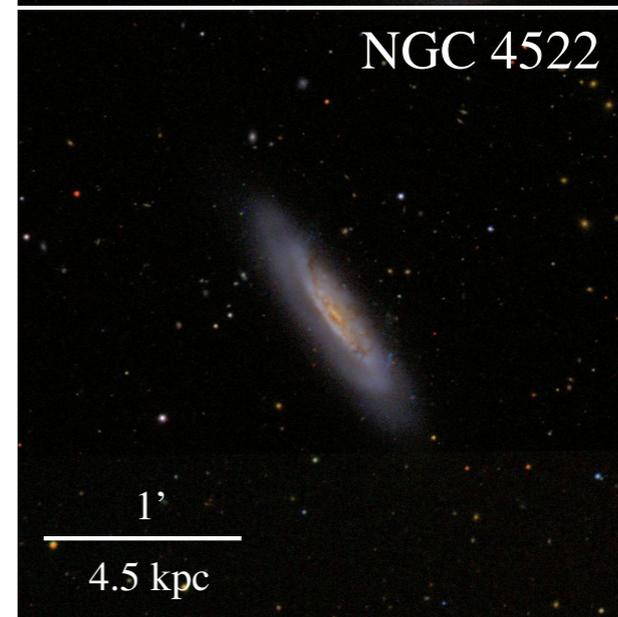
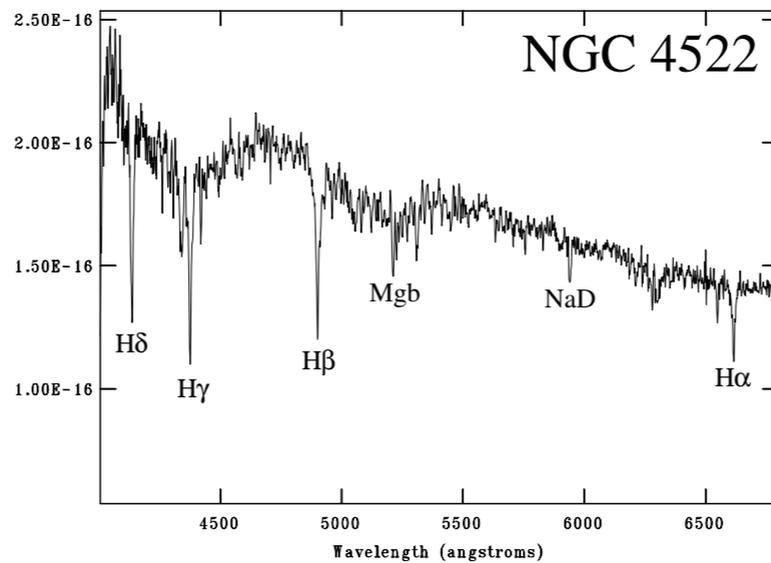
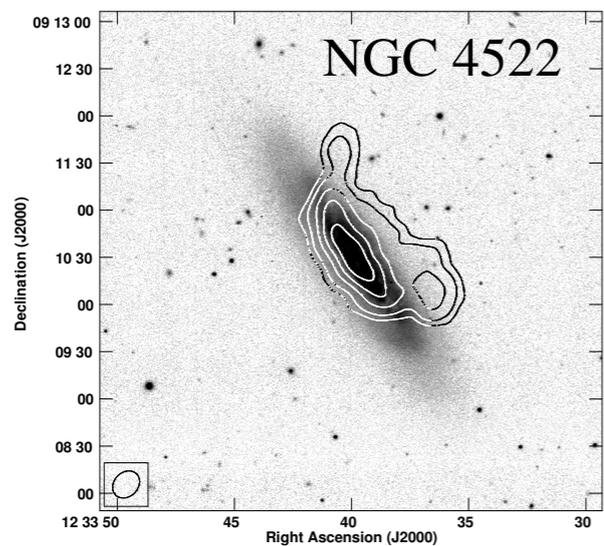
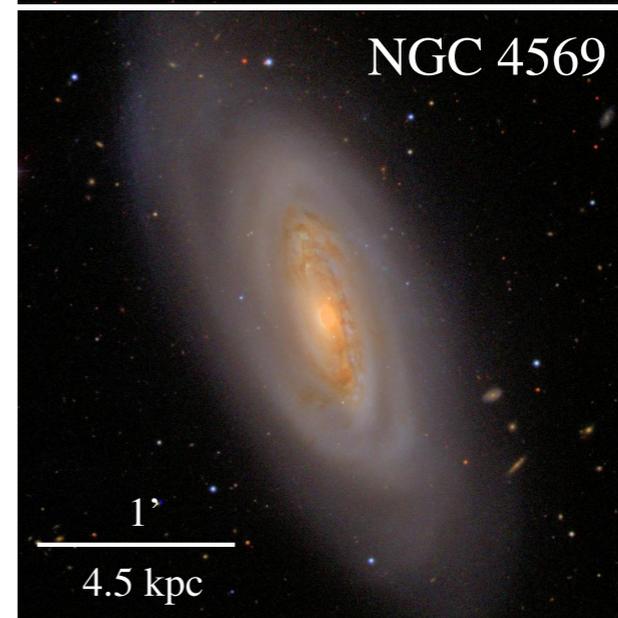
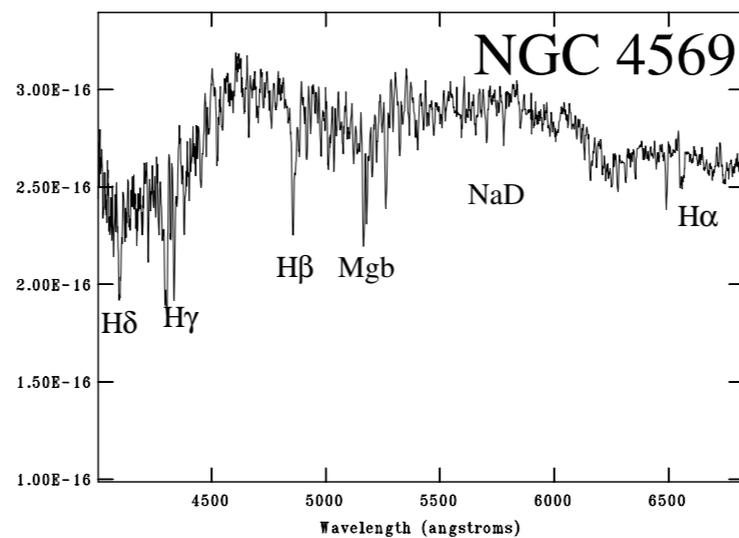
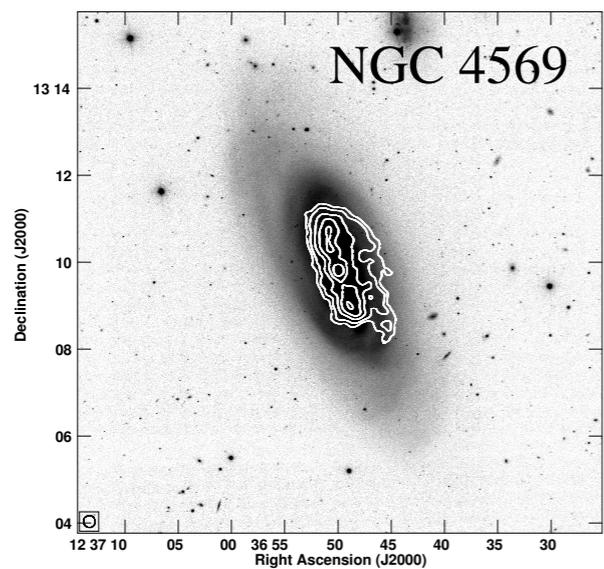
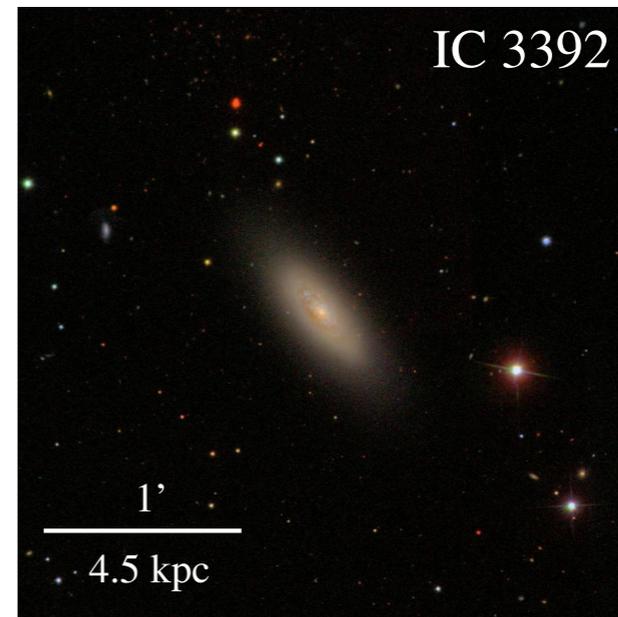
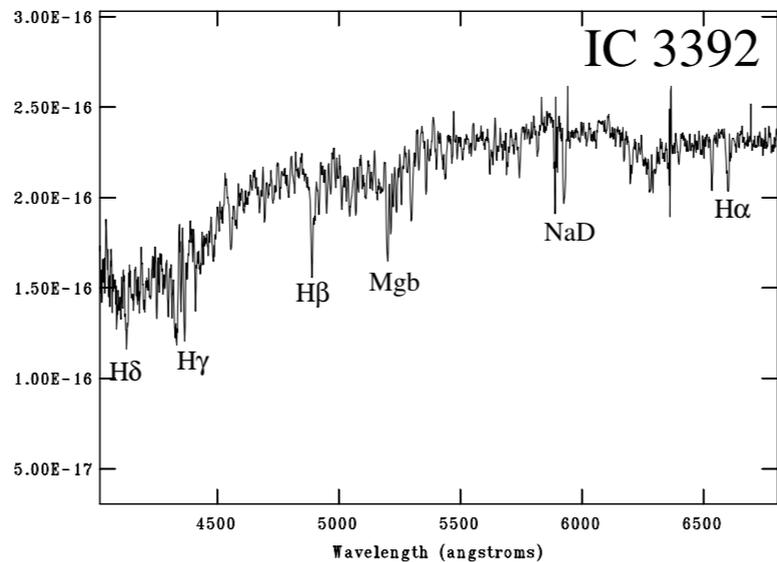
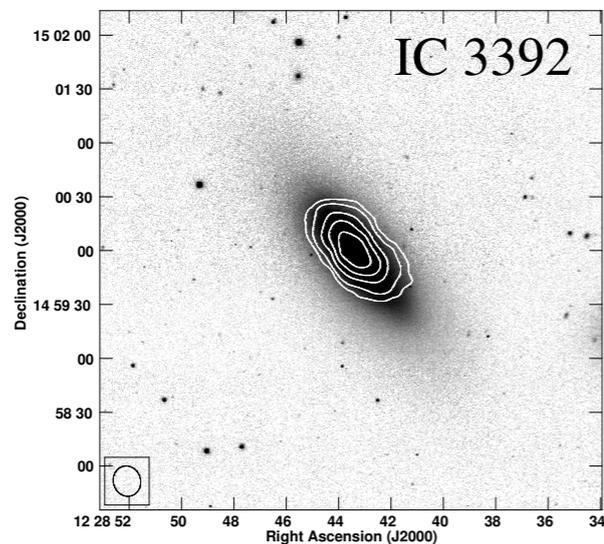
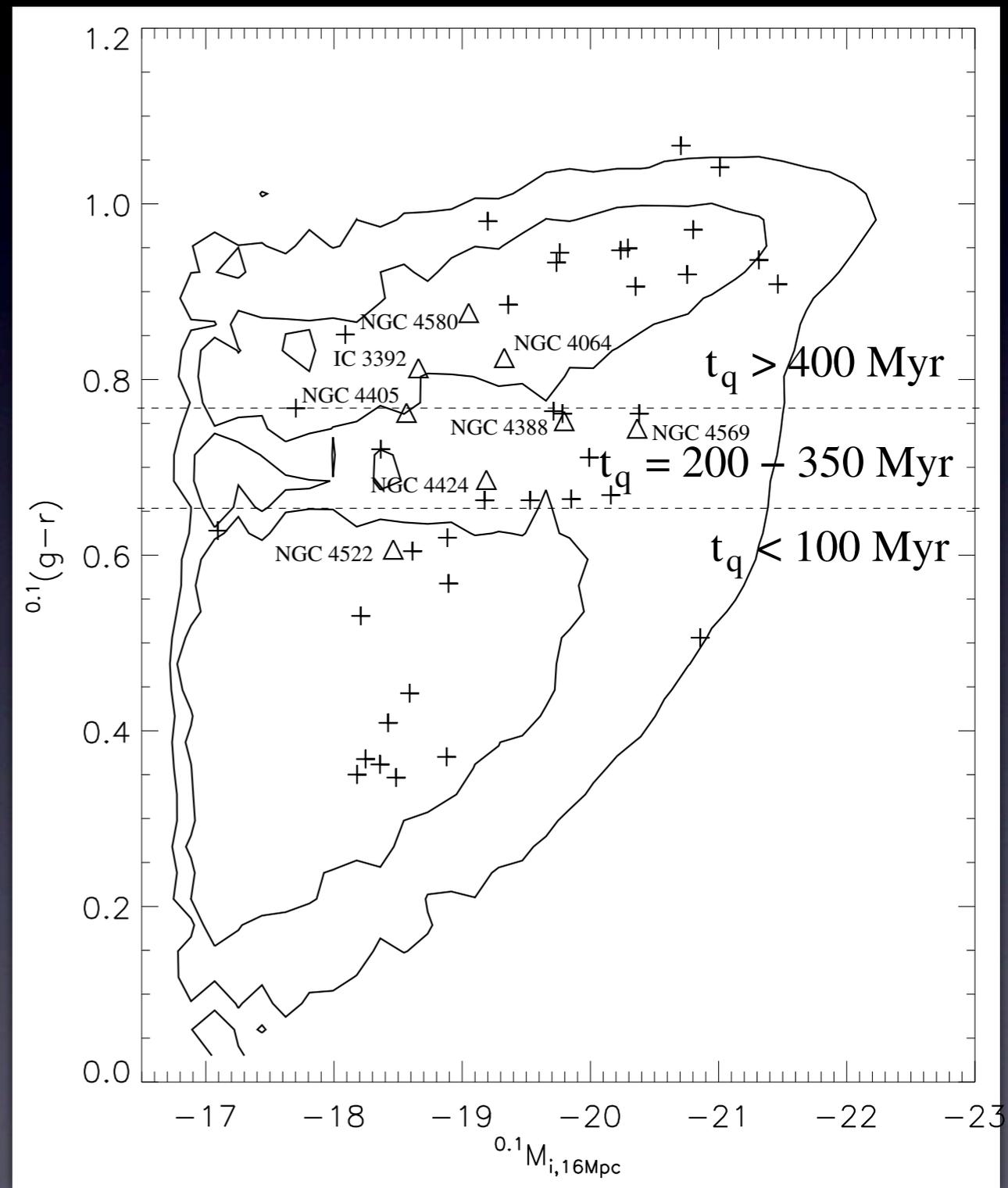
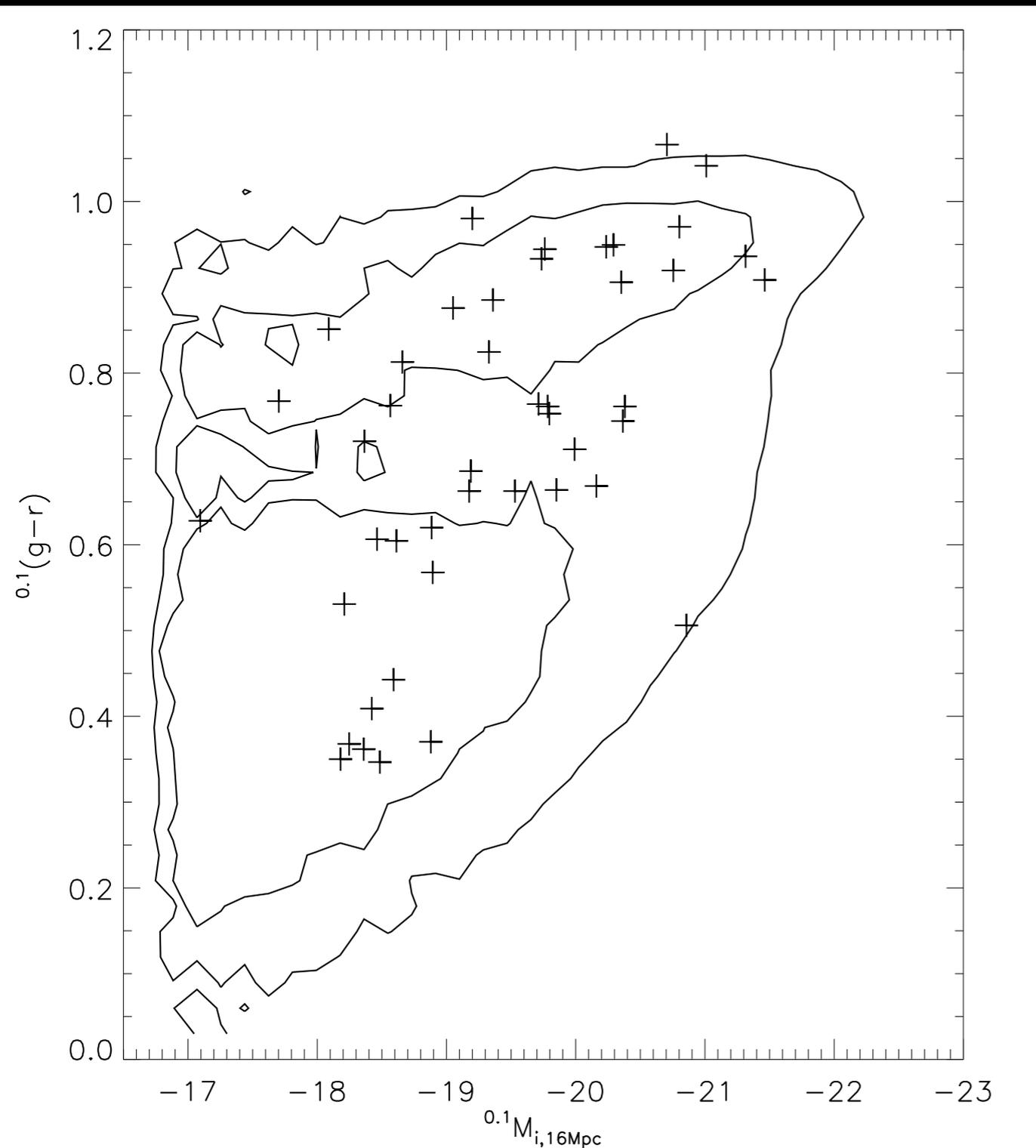


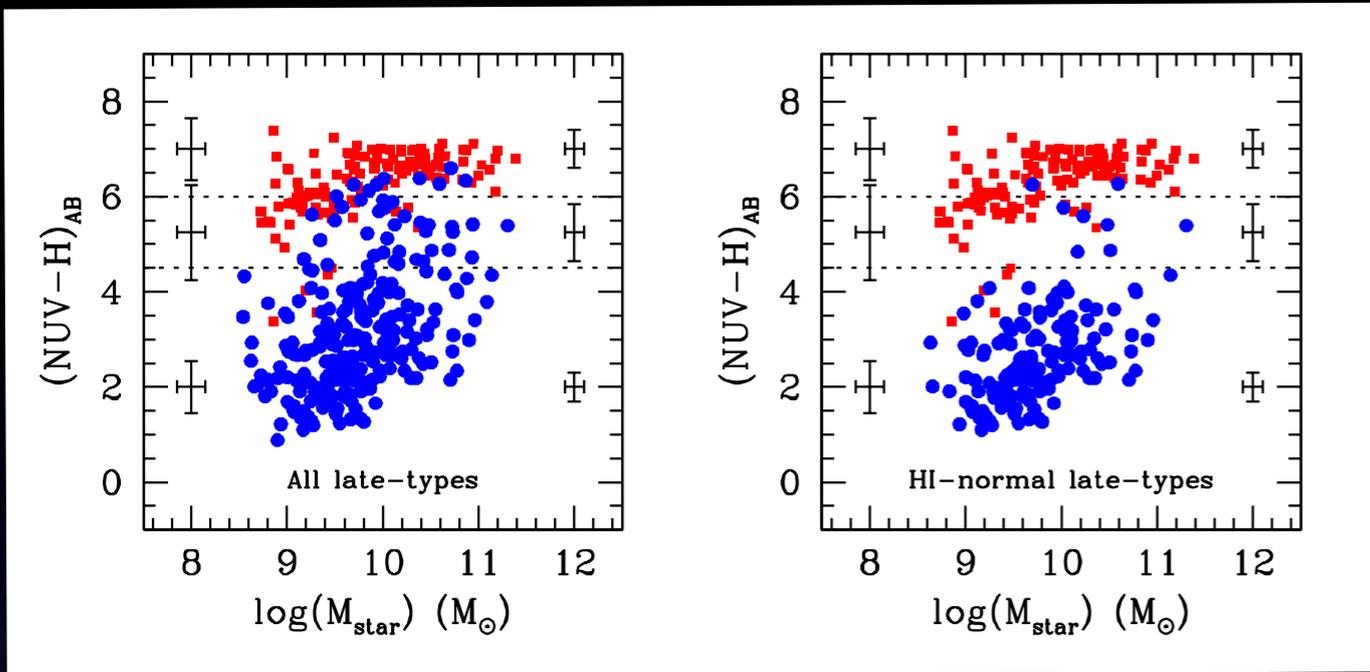
Broadband Colors



Color Magnitude Diagrams



Cortese & Hughes (2009)



HI-normal late types occupy the blue cloud, but HI-deficient galaxies tend to scatter everywhere

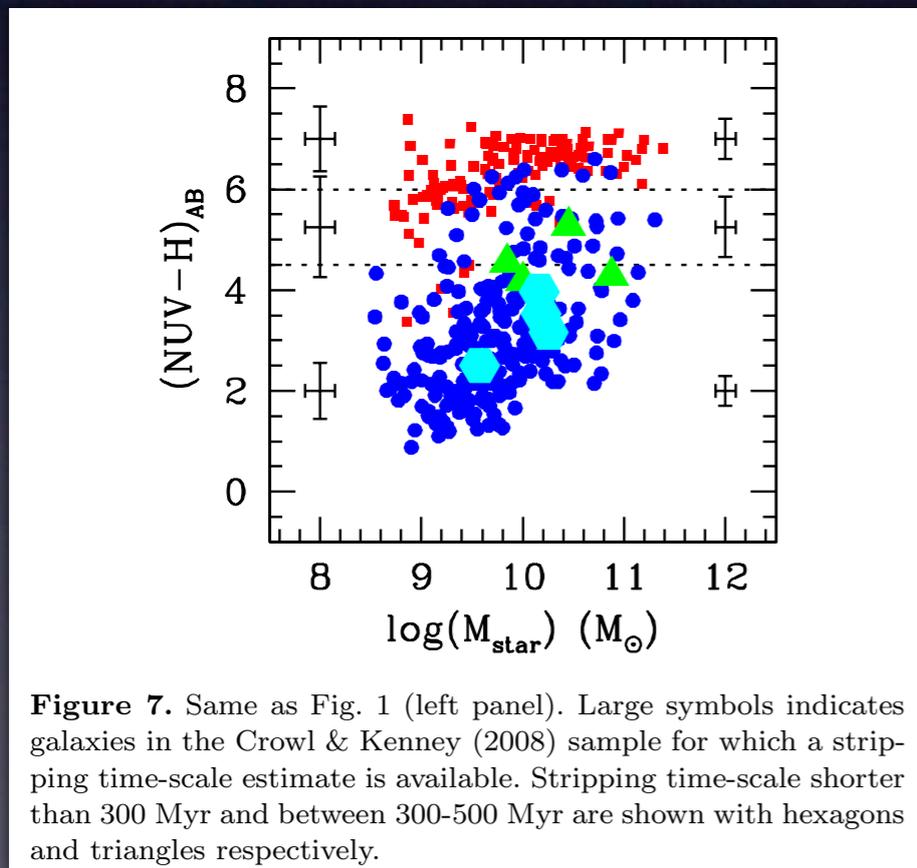


Figure 7. Same as Fig. 1 (left panel). Large symbols indicates galaxies in the Crowl & Kenney (2008) sample for which a stripping time-scale estimate is available. Stripping time-scale shorter than 300 Myr and between 300-500 Myr are shown with hexagons and triangles respectively.

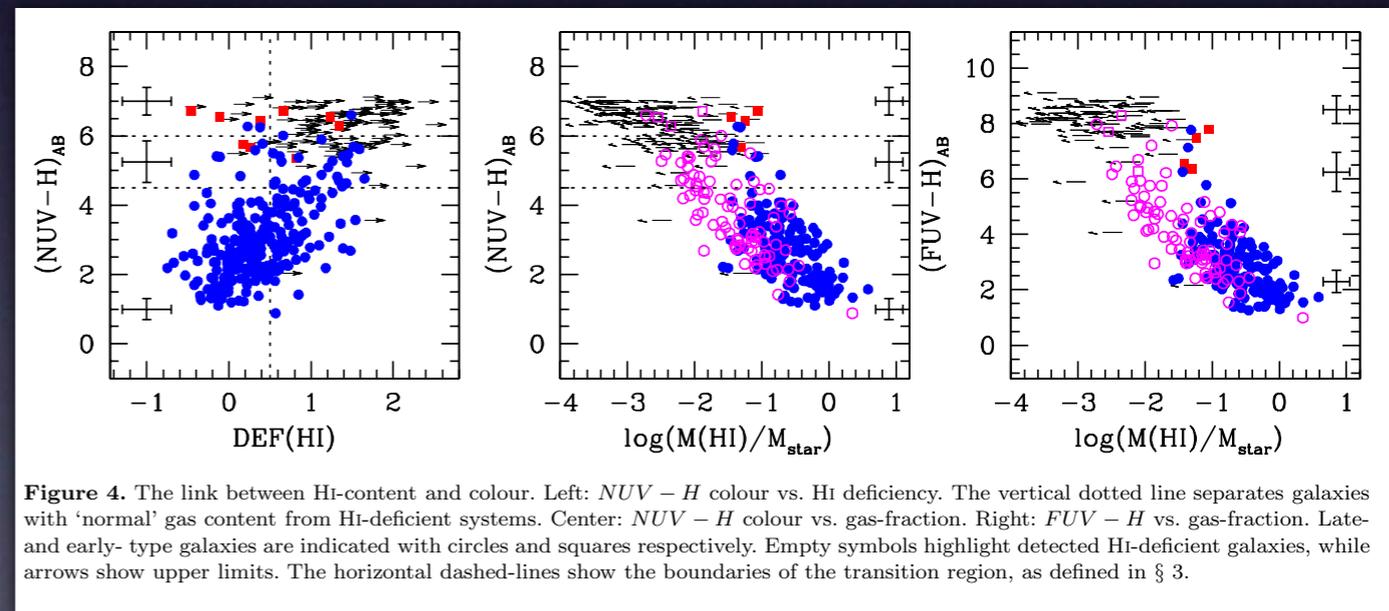
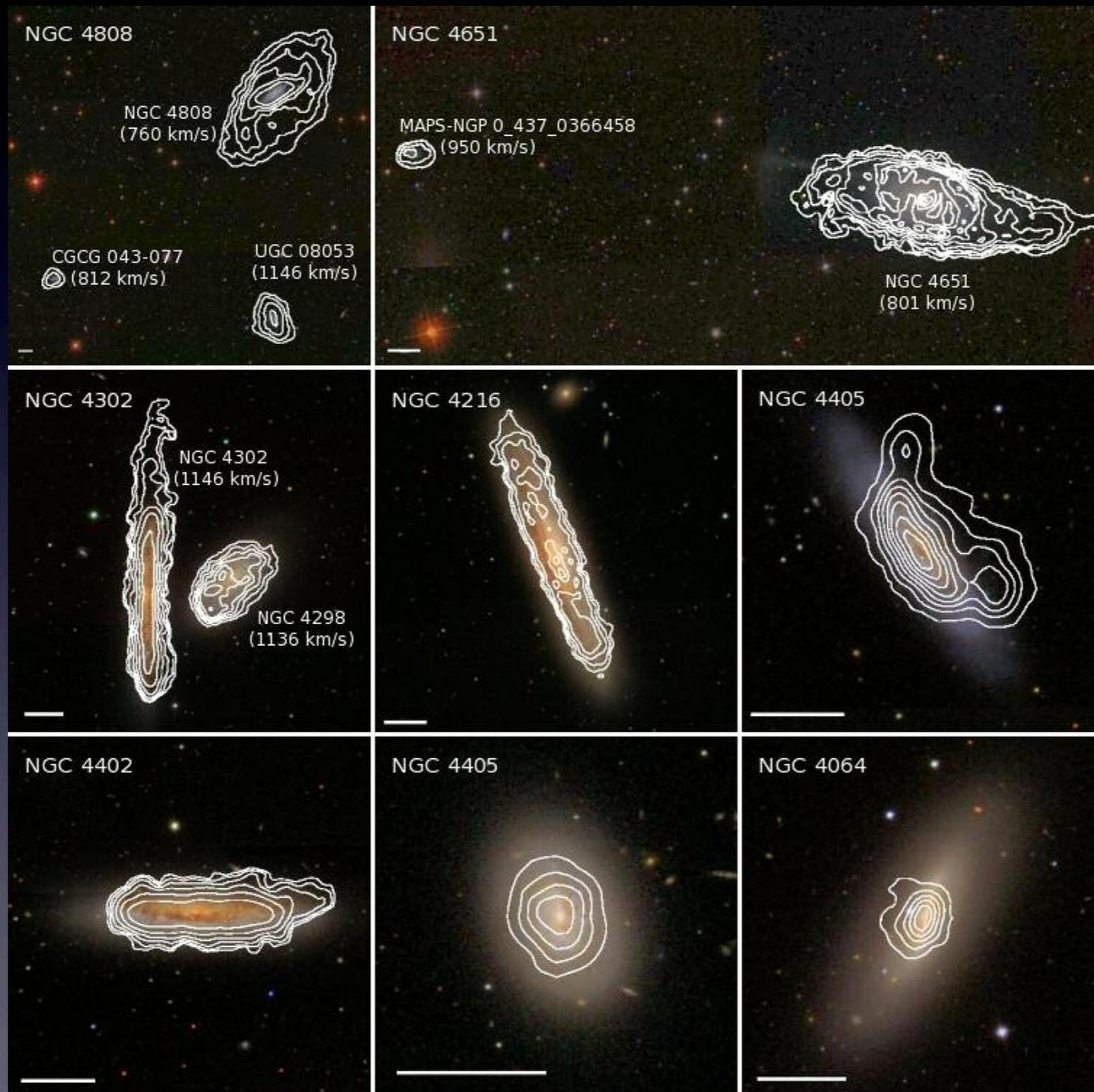


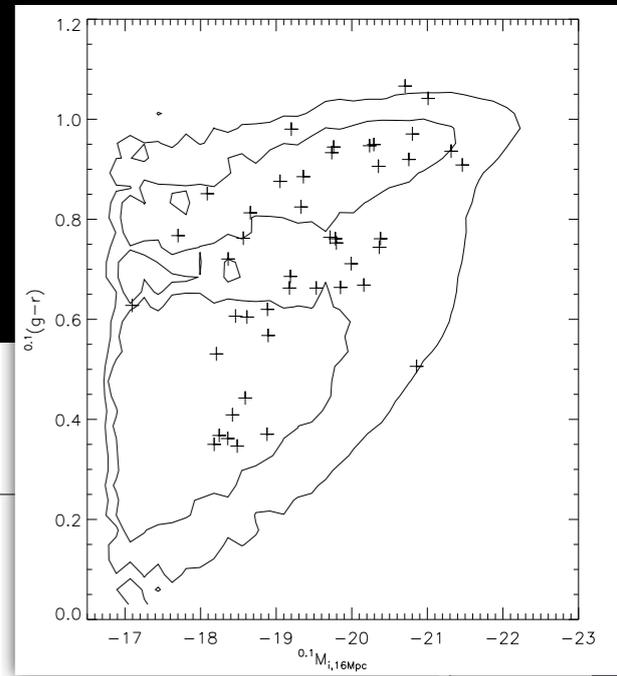
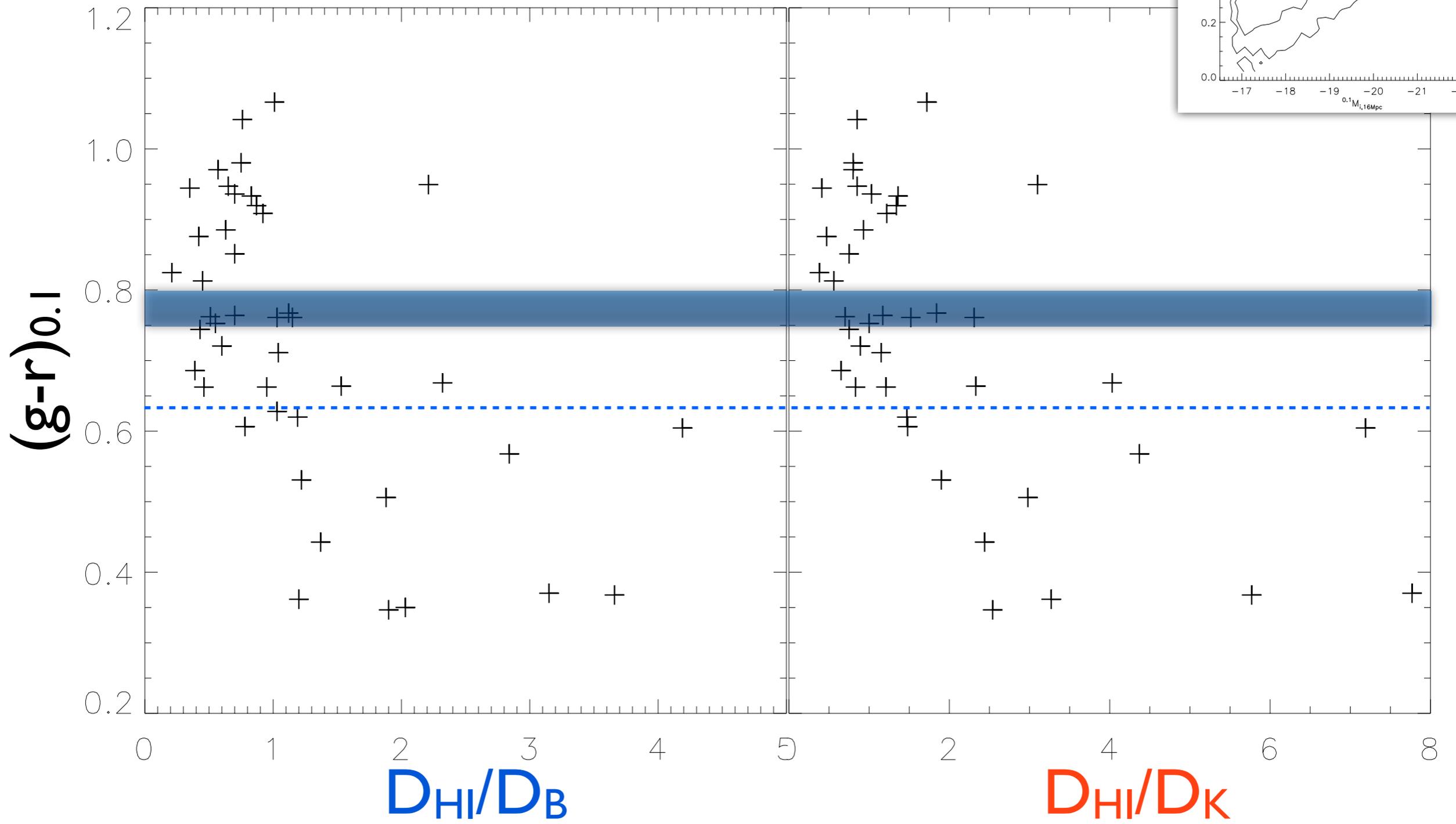
Figure 4. The link between HI-content and colour. Left: $NUV - H$ colour vs. HI deficiency. The vertical dotted line separates galaxies with 'normal' gas content from HI-deficient systems. Center: $NUV - H$ colour vs. gas-fraction. Right: $FUV - H$ vs. gas-fraction. Late- and early- type galaxies are indicated with circles and squares respectively. Empty symbols highlight detected HI-deficient galaxies, while arrows show upper limits. The horizontal dashed-lines show the boundaries of the transition region, as defined in § 3.

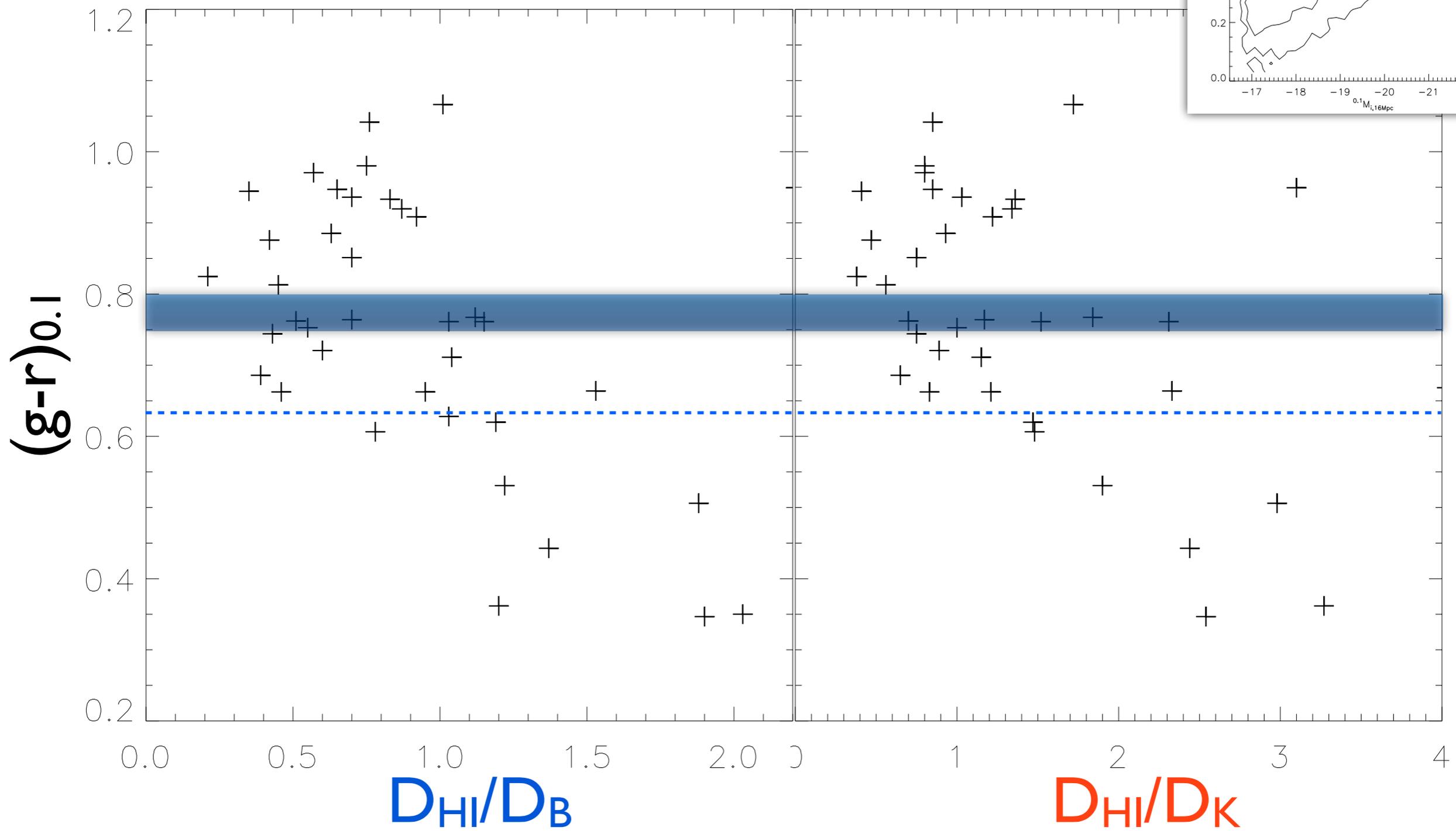
HI-deficiency correlates with color

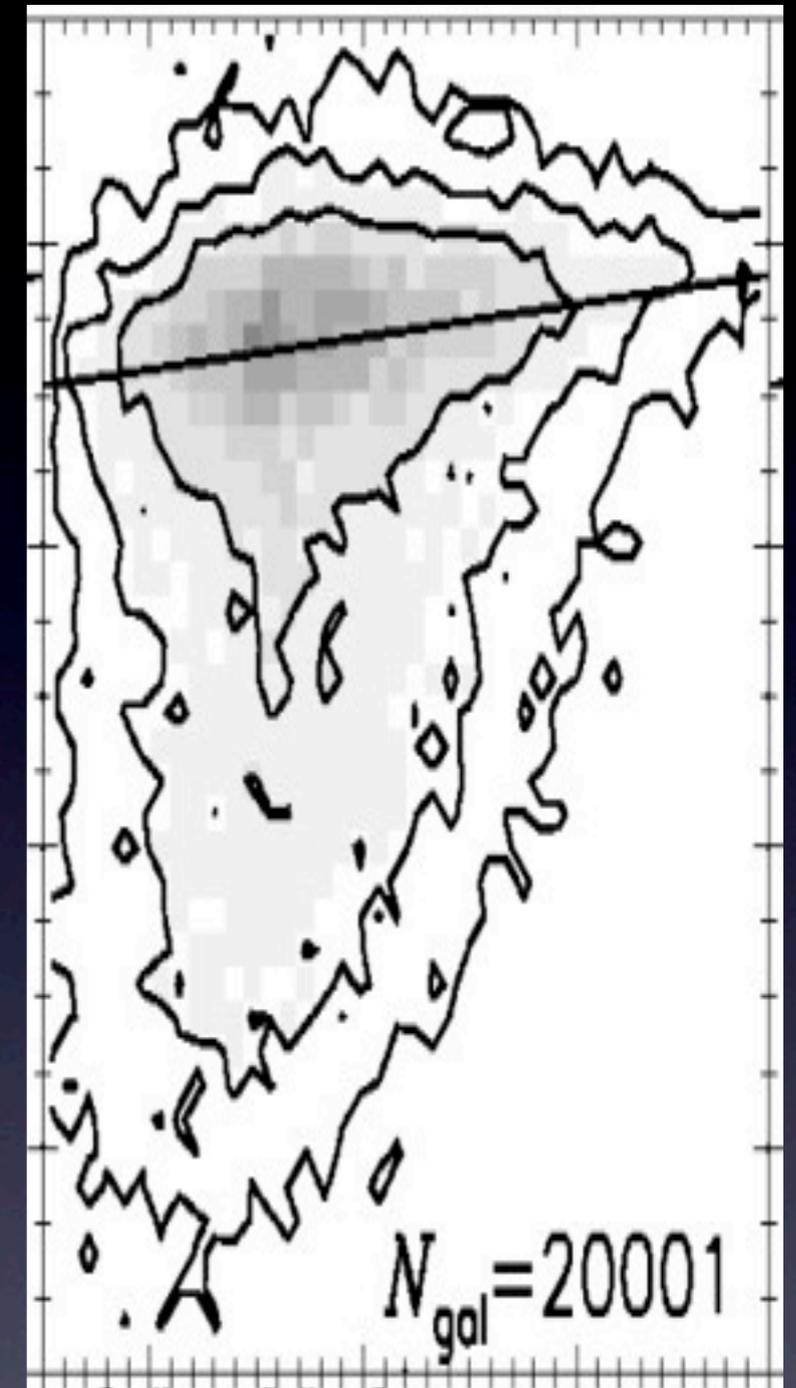
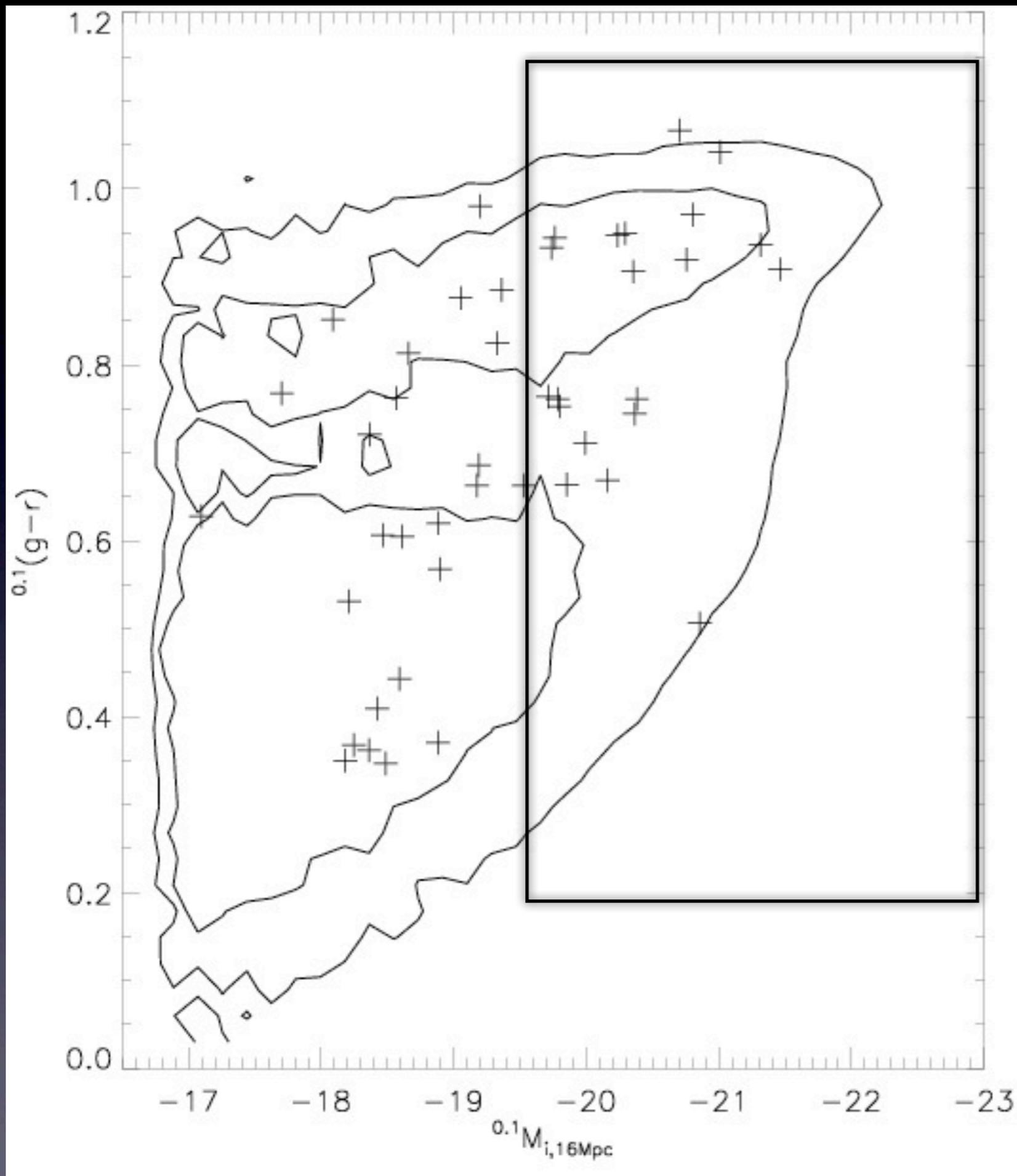
Using our own results against us

But, we know about HI morphologies ...









High density environments
($\delta_{l \times 8} > 7.0$)

All Sersic values

Questions/Thoughts

- **What is the story of the paper?**
 - Galaxies with large HI disks are blue; those with small HI disks are red
 - Galaxies that were stripped very recently occupy the blue cloud, those stripped longer ago occupy the green valley, and those stripped longest ago occupy the red sequence
 - It seems that it takes ~500 Myr to transform from blue to red.
 - Galaxies with disturbed HI disks have very recent star formation and blue colors; those with undisturbed disks have less recent star formation.
- **How are we different from Cortese & Hughes (2009)?**
 - We know HI *morphologies*
 - We know (or think we know) what happened to a lot of these galaxies
- **What should we do about the fact that we are comparing *cluster galaxies to all galaxies?***
 - I think this comparison is important, but it would ALSO be great to compare our sample to a larger “cluster-like” sample.
 - Is it possible to get Blanton (2003)’s CMD with the environment data that Hogg et al. (2004) has for the brightest galaxies?
 - How important is this comparison to *this paper?*

Questions/Thoughts

- Many things different things may have happened to the Virgo Cluster galaxies (RPS, mergers, harassment, starvation (?), etc.)
 - We *only* have timescales for the RPS galaxies; how much can we say about general timescales based on only the RPS results?
 - Is it reasonable to acknowledge that there are many things that happen and then focus on the RPS timescale results
- Other things?