The Galaxy–Dark Matter Connection

constraining cosmology & galaxy formation

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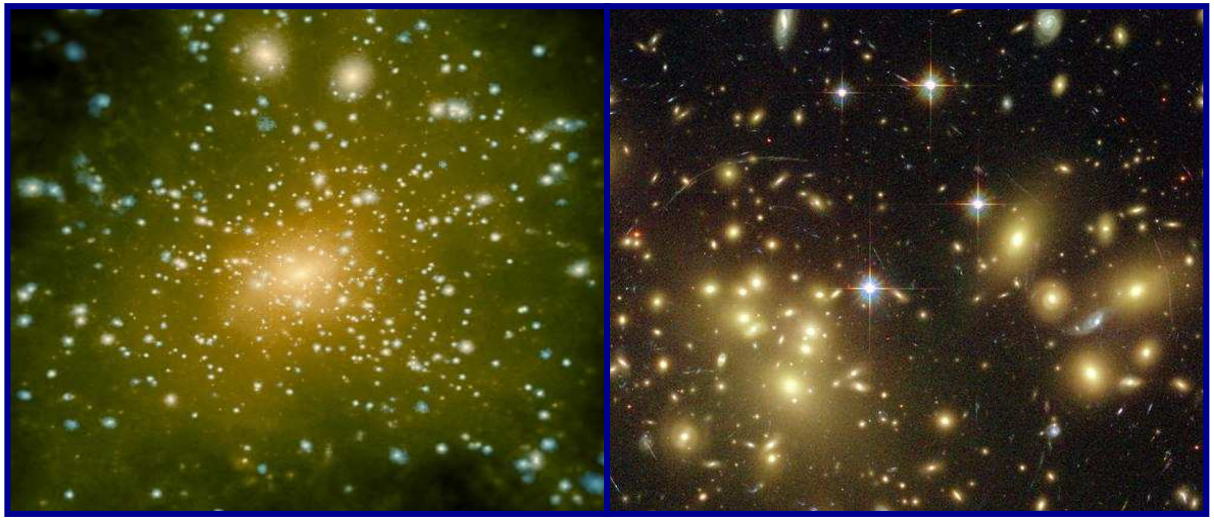
Marcello Cacciato, Surhud More (MPIA)

Outline, Motivation & Techniques

Why study the Galaxy-Dark Matter Connection?

- To constrain the physics of Galaxy Formation
- To constrain Cosmological Parameters



Four Methods to Constrain Galaxy-Dark Matter Connection

- Group Catalogues
- Satellite Kinematics
- Large Scale Structure
- Galaxy-Galaxy Lensing

New Cosmological Constraints

- Precision cosmology using non-linear structure

The Conditional Luminosity Function

In order to parameterize the Halo Occupation Statistics we introduce the Conditional Luminosity Function (CLF), $\Phi(L|M)$, which is the direct link between the halo mass function $n(M)$ and the galaxy luminosity function $\Phi(L)$:

$$\Phi(L) = \int_0^\infty \Phi(L|M)\, n(M)\, \mathrm{d}M$$

The CLF contains a wealth of information, such as:

- **The average relation between light and mass:**

$$\langle L\rangle(M) = \int_0^\infty \Phi(L|M)\, L\, \mathrm{d}L$$

- **The occupation numbers of galaxies:**

$$\langle N\rangle(M) = \int_{L_{\rm min}}^\infty \Phi(L|M)\, \mathrm{d}L$$

We constrain CLF using four different, independent techniques

Galaxy Group Catalogues | **Large Scale Structure**
Satellite Kinematics | **Galaxy-Galaxy Lensing**

The CLF Model

We split CLF in central and satellite components

$$\Phi(L|M)\mathrm{d}L = \Phi_{\mathrm{c}}(L|M)\mathrm{d}L + \Phi_{\mathrm{s}}(L|M)\mathrm{d}L$$

- **For centrals we adopt a log-normal distribution**

$$\Phi_{\mathrm{c}}(L|M)\mathrm{d}L = \frac{1}{\sqrt{2\pi}\,\sigma_{\mathrm{c}}} \exp\left[-\left(\frac{\ln(L/L_{\mathrm{c}})}{\sqrt{2}\sigma_{\mathrm{c}}}\right)^2\right] \frac{\mathrm{d}L}{L}$$

- **For satellites we adopt a modified Schechter function**

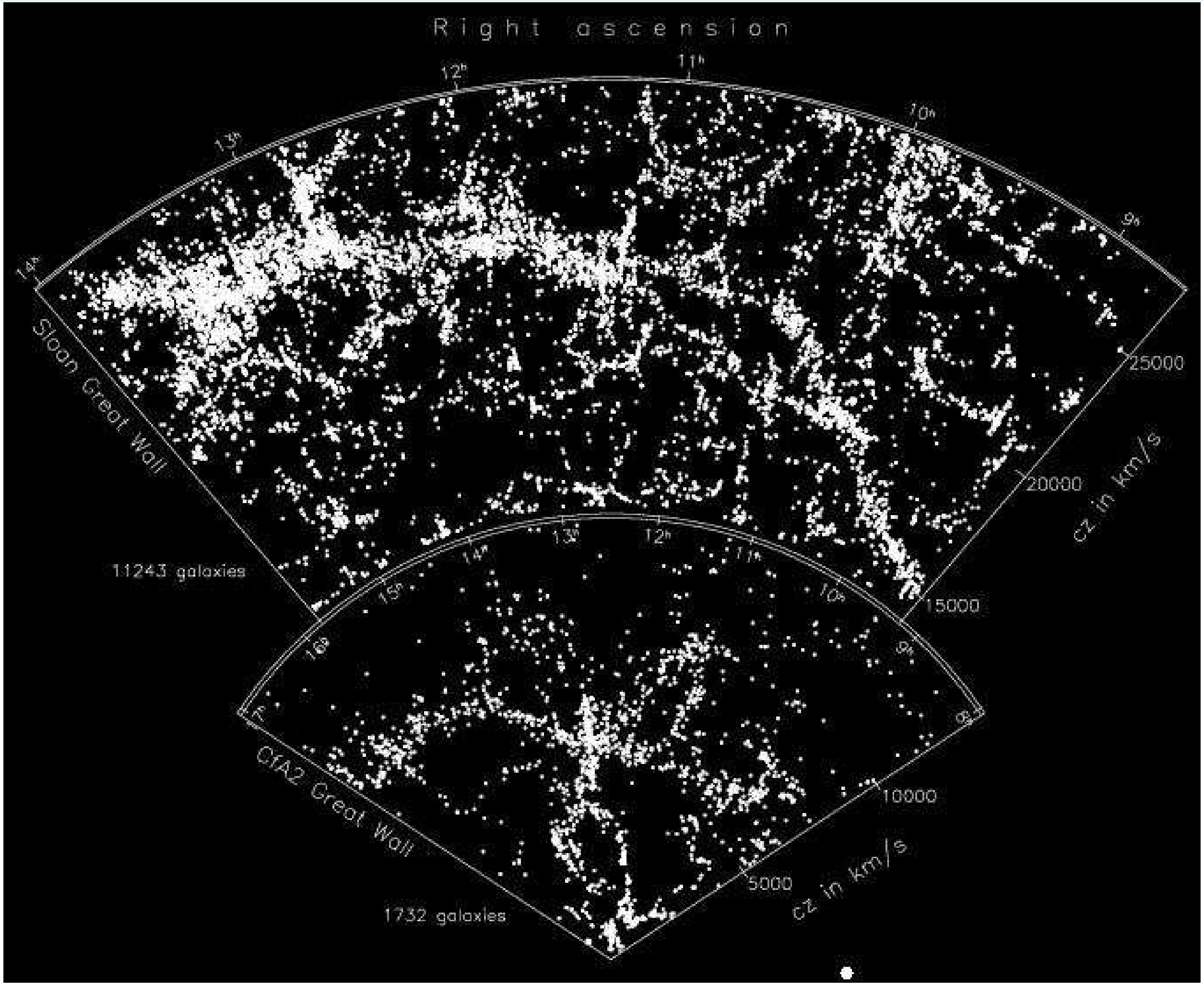
$$\Phi_{\mathrm{s}}(L|M)\mathrm{d}L = \frac{\Phi_{\mathrm{s}}}{L_{\mathrm{s}}} \left(\frac{L}{L_{\mathrm{s}}}\right)^{\alpha_{\mathrm{s}}} \exp[-(L/L_{\mathrm{s}})^2]\,\mathrm{d}L$$

Note that L_{c}, L_{s}, σ_{c}, ϕ_{s} and α_{s} all depend on halo mass M

Free parameters are constrained by fitting data

Use Monte-Carlo Markov Chain to sample the posterior distribution of free parameters, and to put confidence levels on derived quantities

Galaxy Groups from Redshift Surveys



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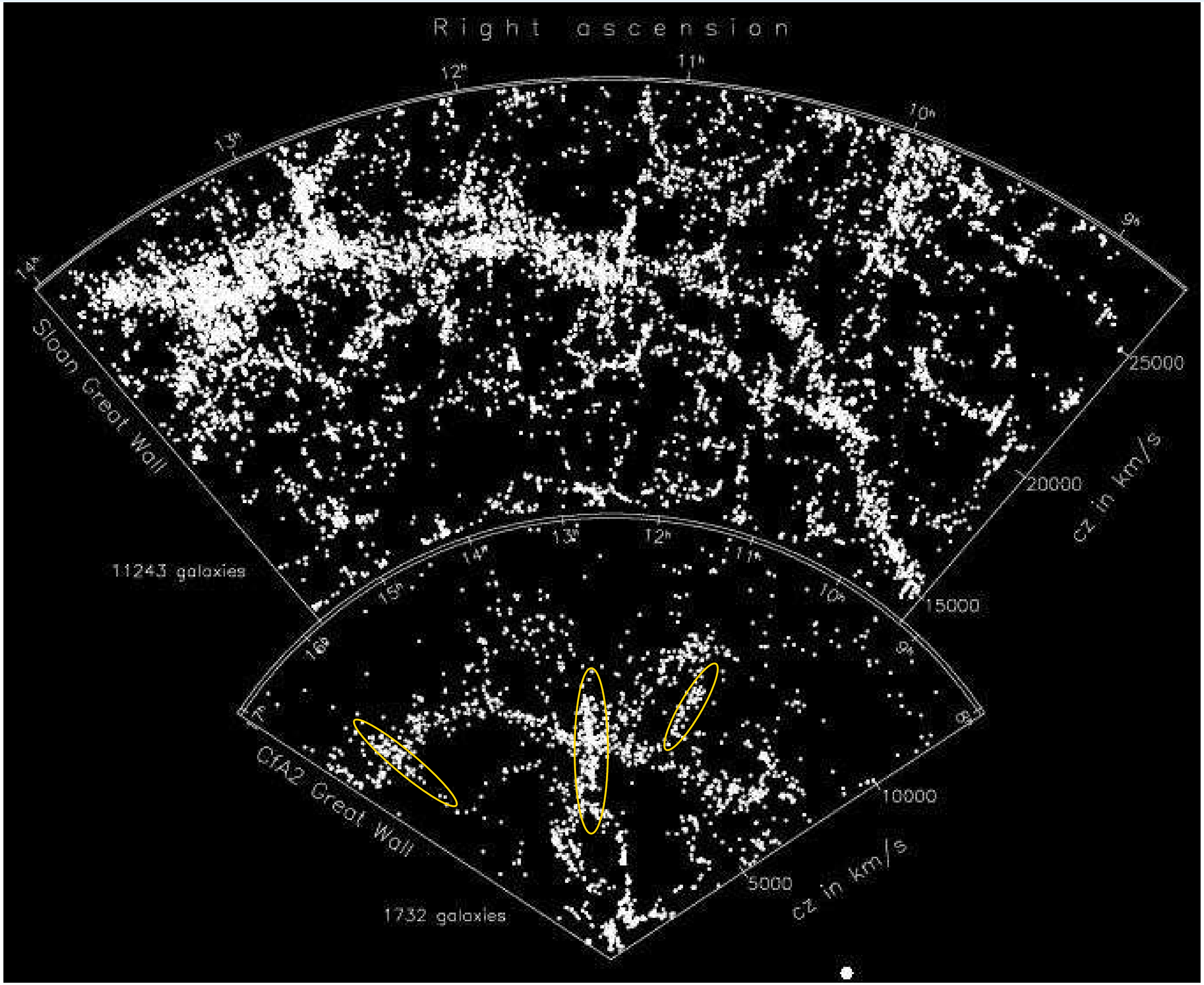

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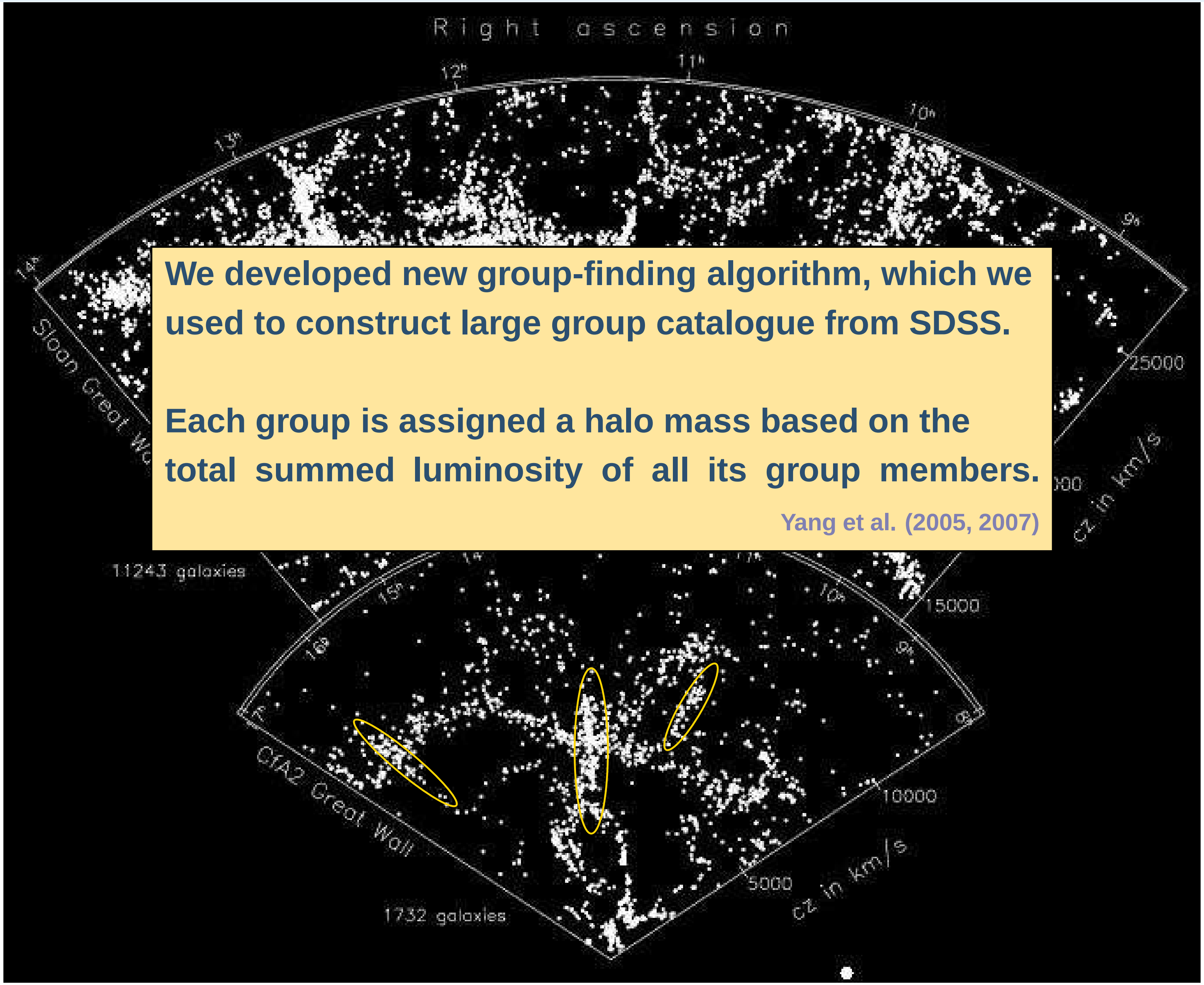
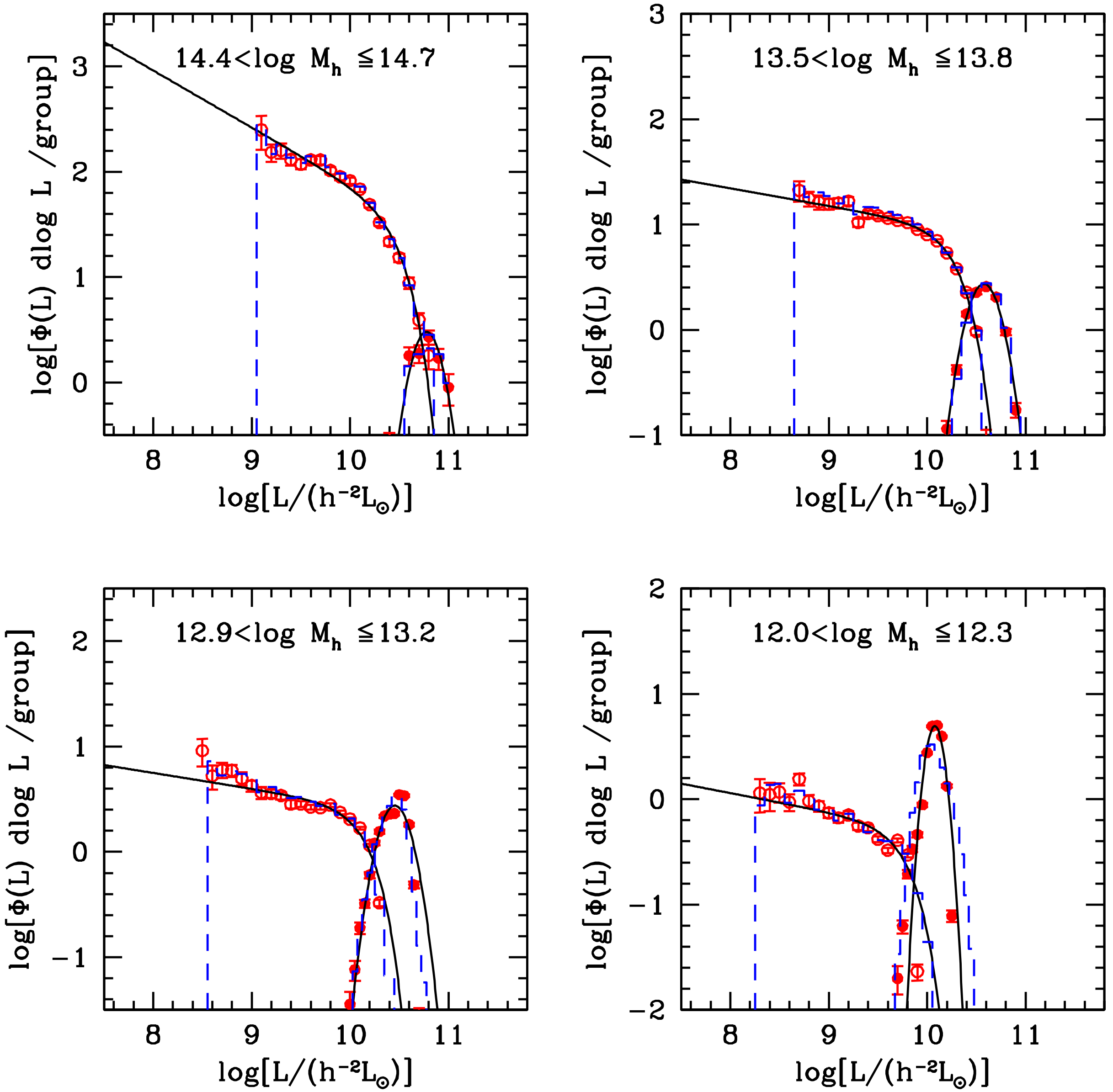


We developed new group-finding algorithm, which we used to construct large group catalogue from SDSS.

Each group is assigned a halo mass based on the total summed luminosity of all its group members.

Yang et al. (2005, 2007)

The CLF from SDSS Group Catalogue



Yang, Mo, vdB (2007)

Occupation Statistics from Clustering

- Galaxies occupy dark matter halos.
- CDM: more massive halos are more strongly clustered.
- Clustering strength of given population of galaxies indicates the characteristic halo mass

Clustering strength measured by correlation length r_0

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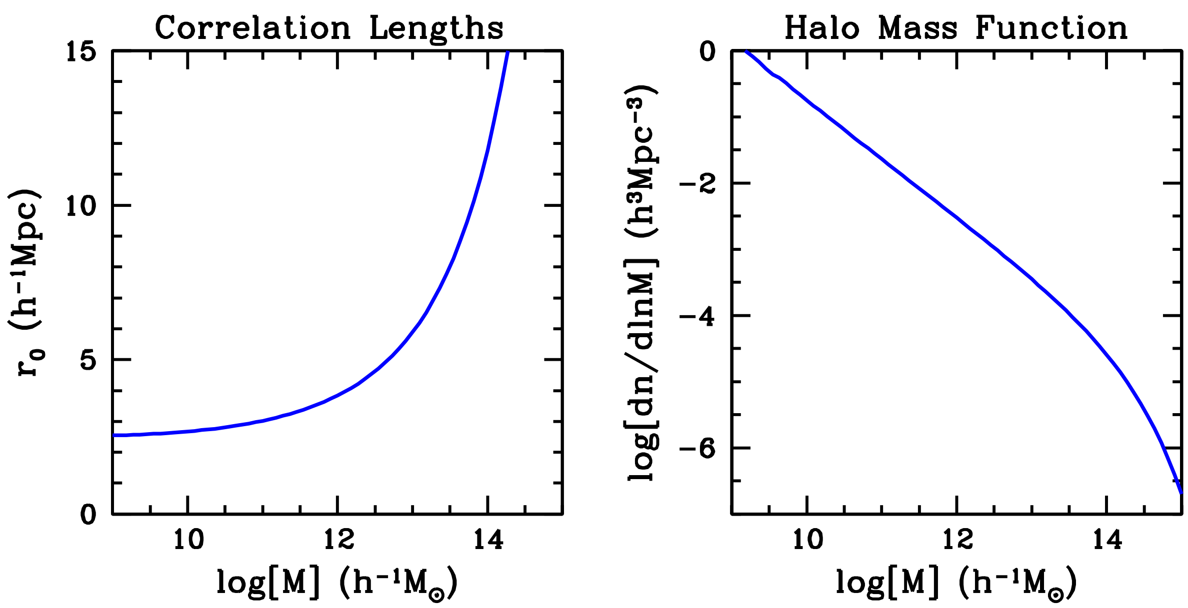

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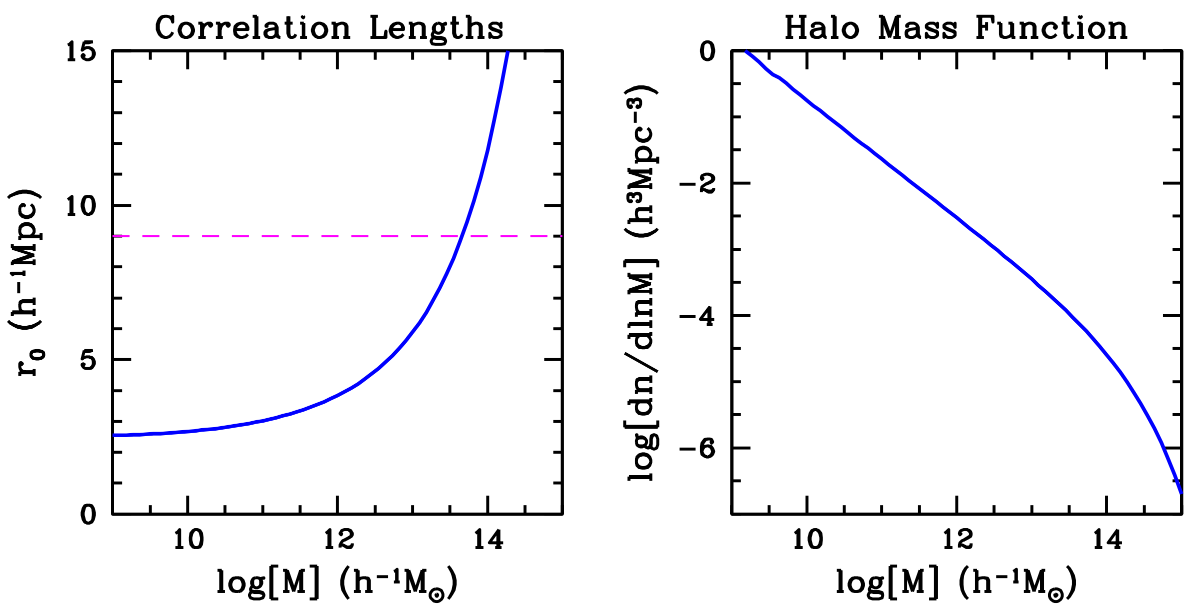

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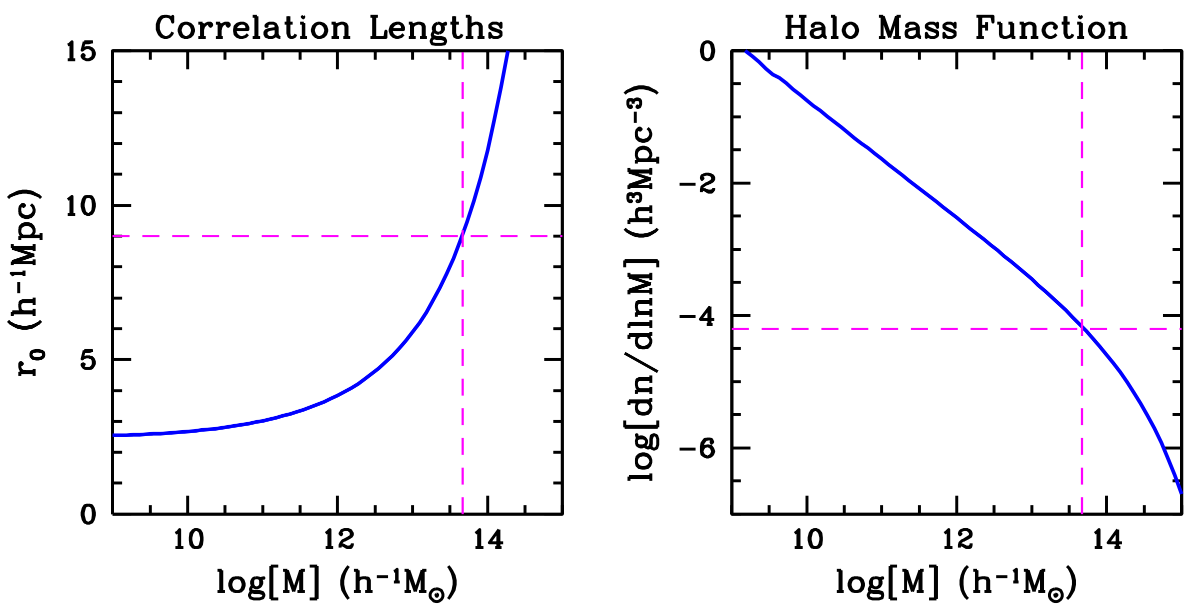

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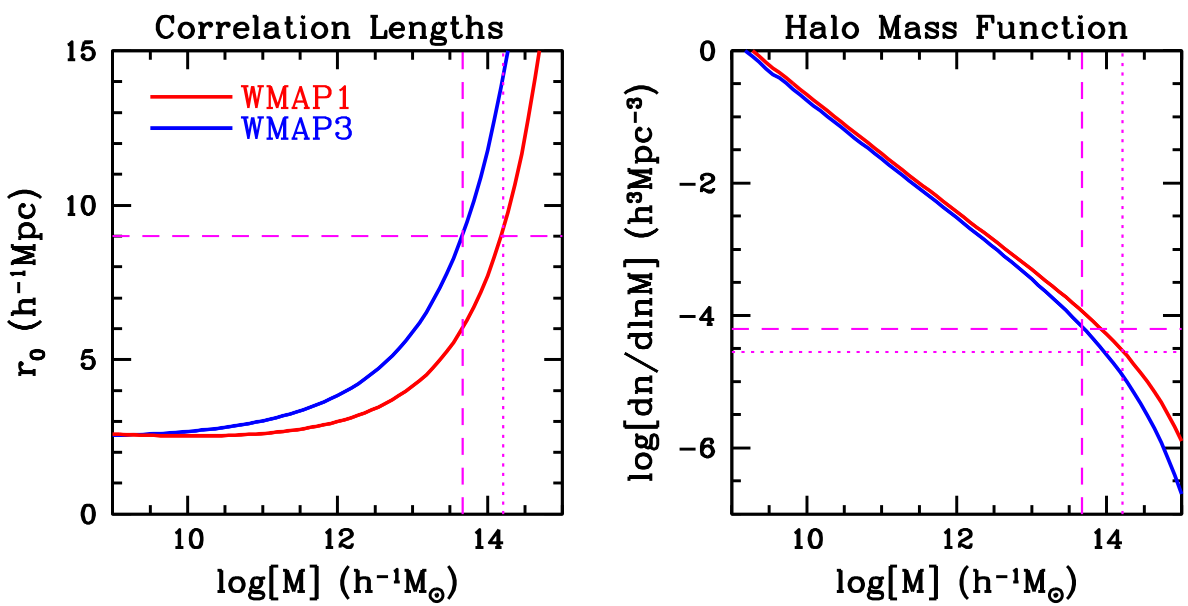


WMAP1

$\Omega_m = 0.30$

$\Omega_\Lambda = 0.70$

$\sigma_8 = 0.90$

WMAP3

$\Omega_m = 0.24$

$\Omega_\Lambda = 0.76$

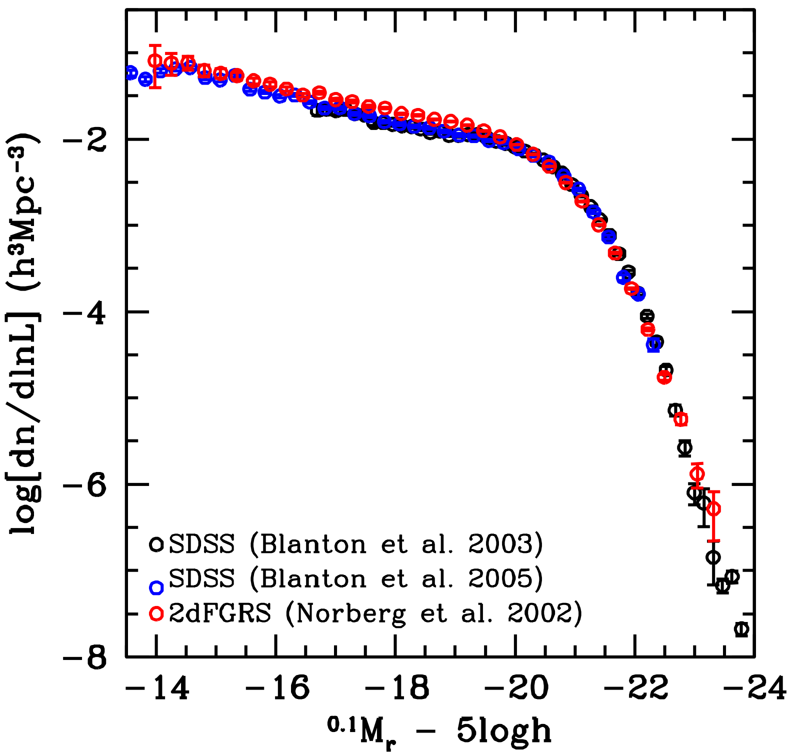
$\sigma_8 = 0.74$

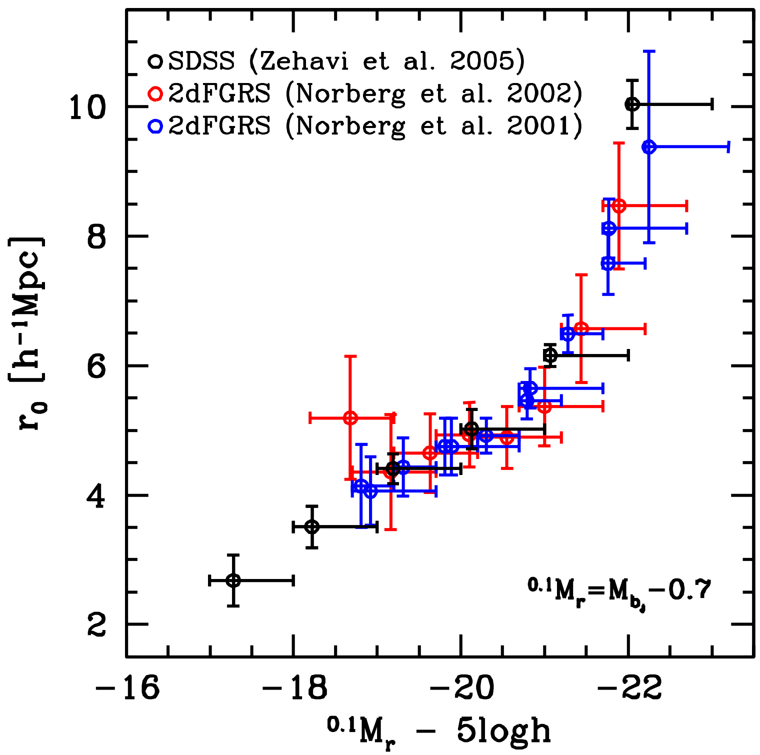
CAUTION: Results depend on cosmological parameters

Luminosity & Correlation Functions

Introduction
Conditional Luminosity Function
Galaxy Group Catalogues
Large Scale Structure
- Occupation Statistics from Clustering
- Luminosity & Correlation Functions
- Results
- Cosmology Dependence

Satellite Kinematics
Galaxy-Galaxy Lensing
Conclusions
Extra Material



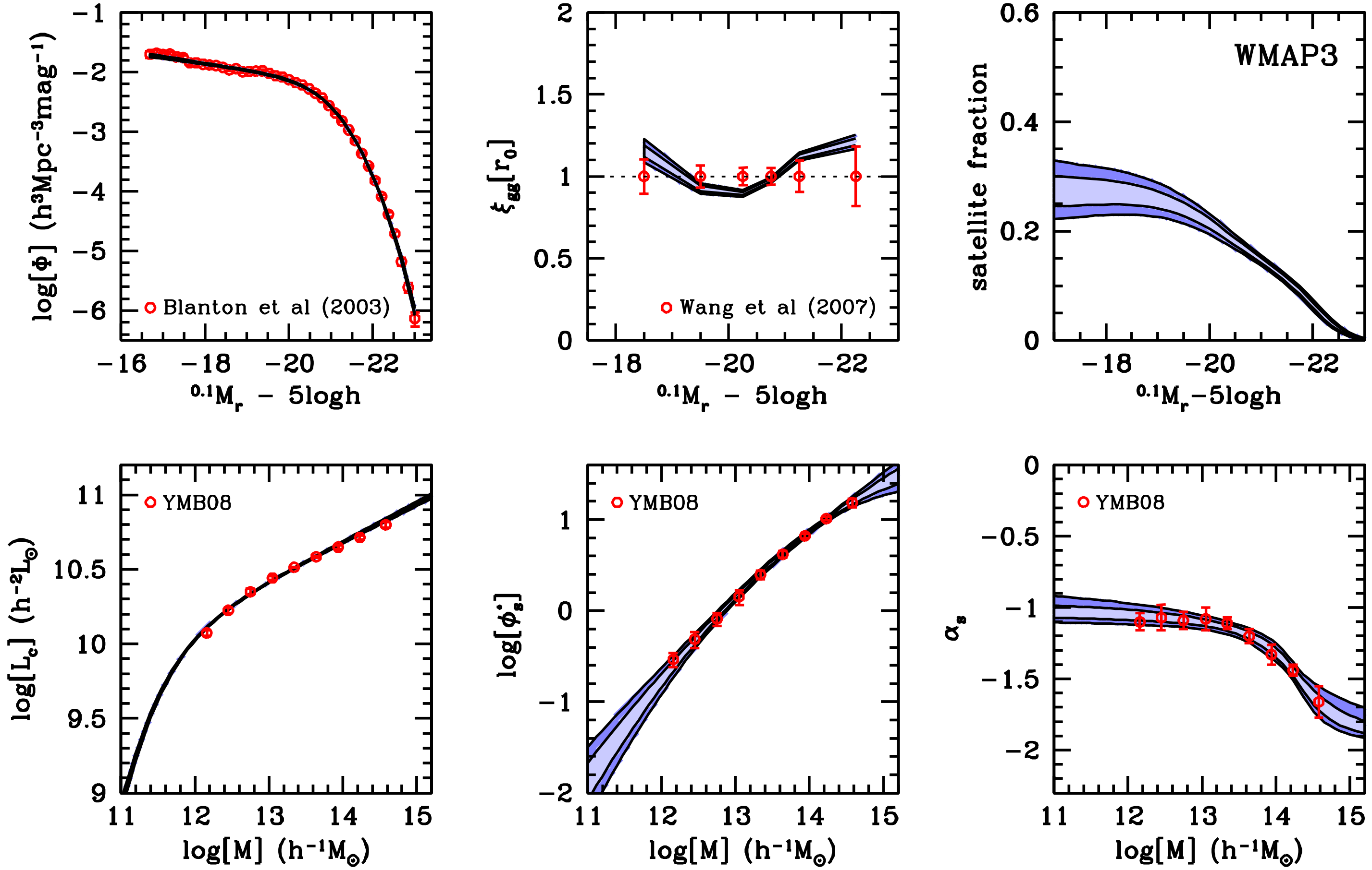


- **DATA:** More luminous galaxies are more strongly clustered.
- **ΛCDM:** More massive haloes are more strongly clustered.

More luminous galaxies reside in more massive haloes

REMINDER: Correlation length r_0 defined by $\xi(r_0) = 1$

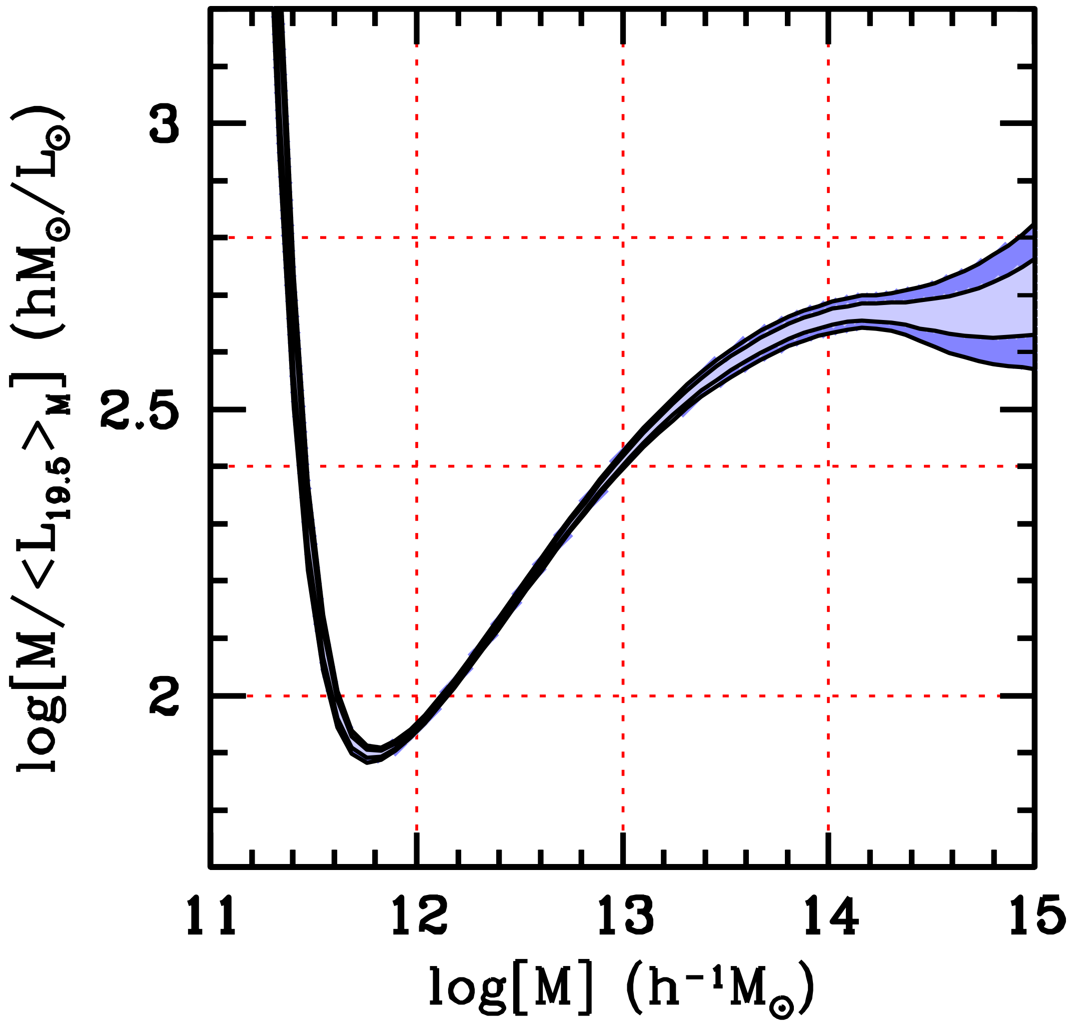
Results



(Cacciato, vdB et al. 2008)

Model fits data extremely well with $\chi^2_{\rm red} \sim 1$
Same model in excellent agreement with results from SDSS galaxy group catalogue of Yang et al. (2008)

Cosmology Dependence



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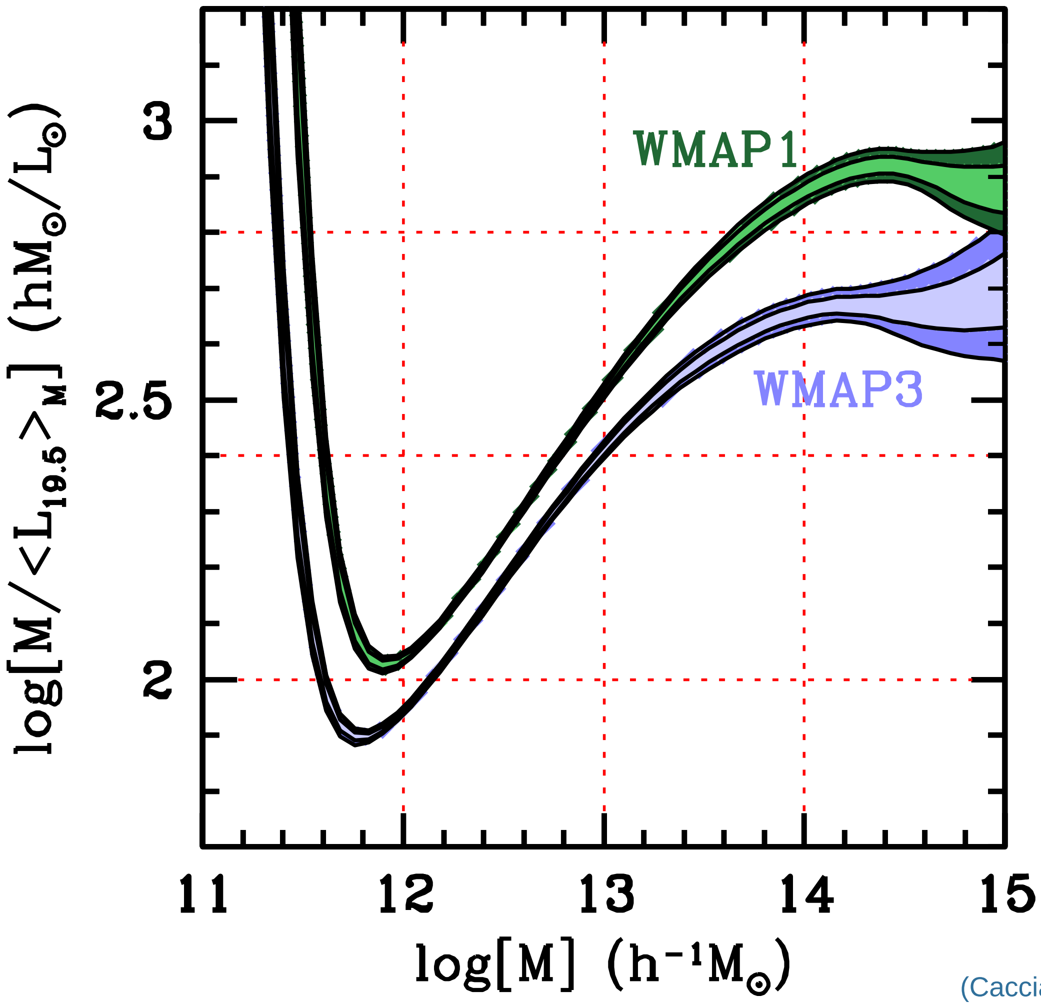


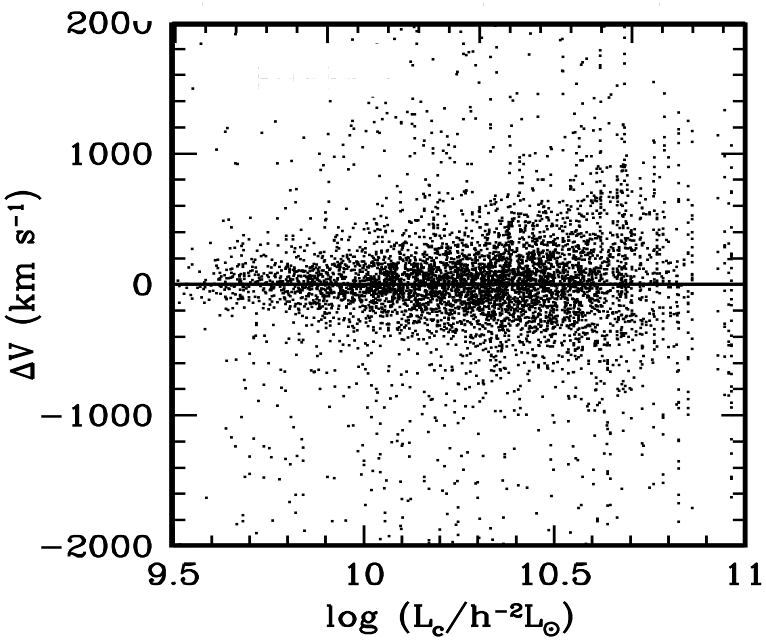
(Cacciato, vdB et al. 2008)

Mass-to-Light ratios tightly constrained, but with strong dependence on cosmology

Satellite Kinematics: Methodology

Select centrals and their satellites from a redshift survey

Using redshifts, determine $\Delta V = V_{\rm sat} - V_{\rm cen}$ as function of L_c



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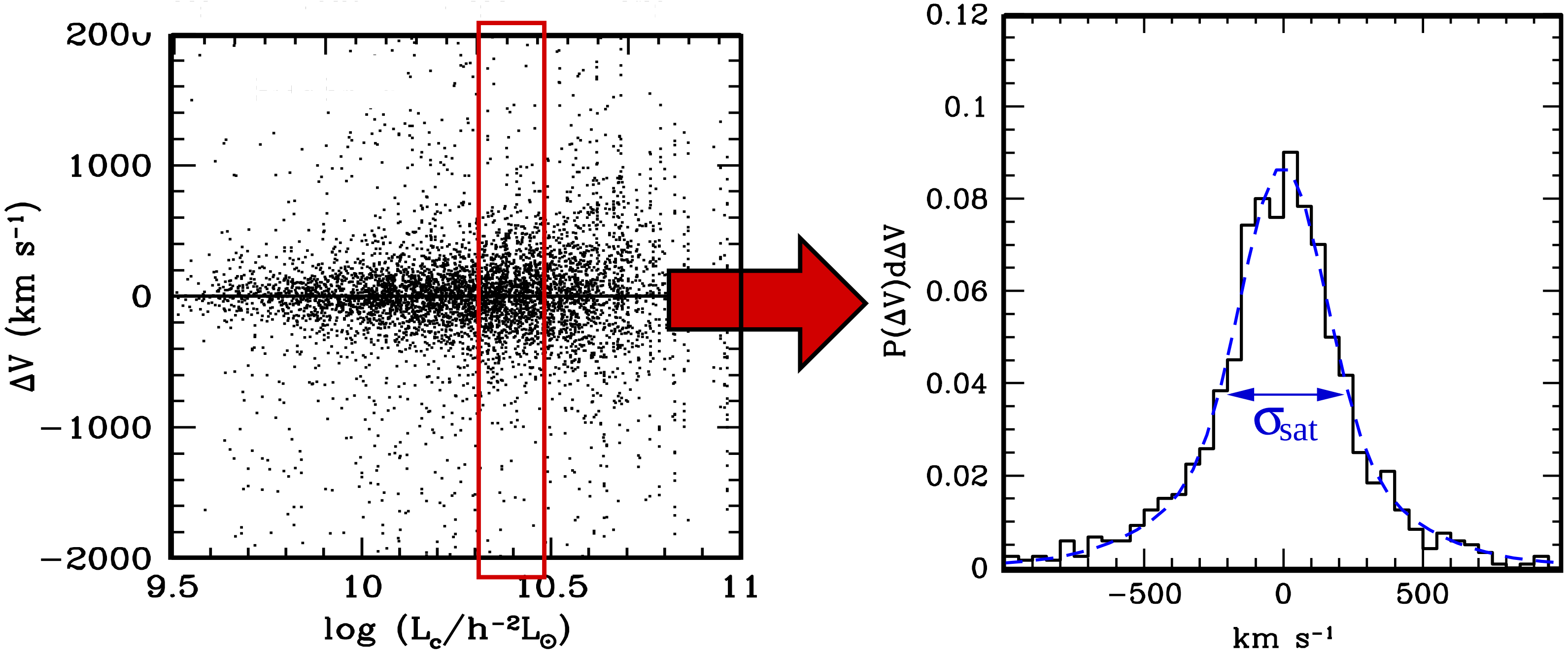

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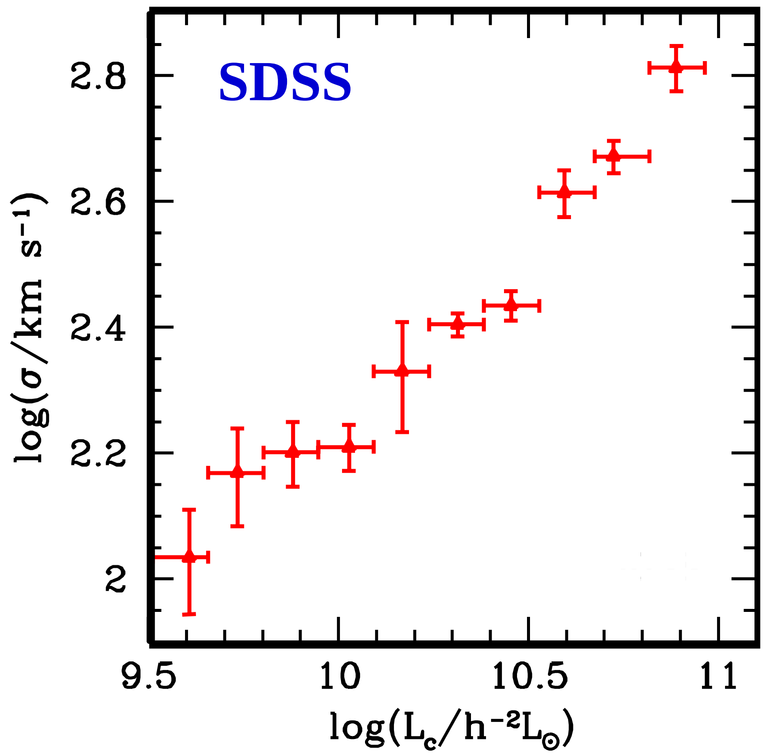

(More, vdB et al. 2008)

Brighter centrals reside in more massive haloes.

Satellite Kinematics: Mass Estimates

Using virial equilibrium and spherical collapse model:

$$\sigma^2 \propto \frac{GM}{R} \qquad M \propto R^3 \qquad \sigma \propto M^{1/3}$$

Satellite Kinematics: Mass Estimates

Using **virial equilibrium** and **spherical collapse model**:

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On average only ~ 2 satellites per central $\longrightarrow$ **stacking**

Unless $P(M|L_c)$ is a Dirac delta function, stacking means combining halos of different masses

Consequently, one has to distinguish two different weighting schemes:

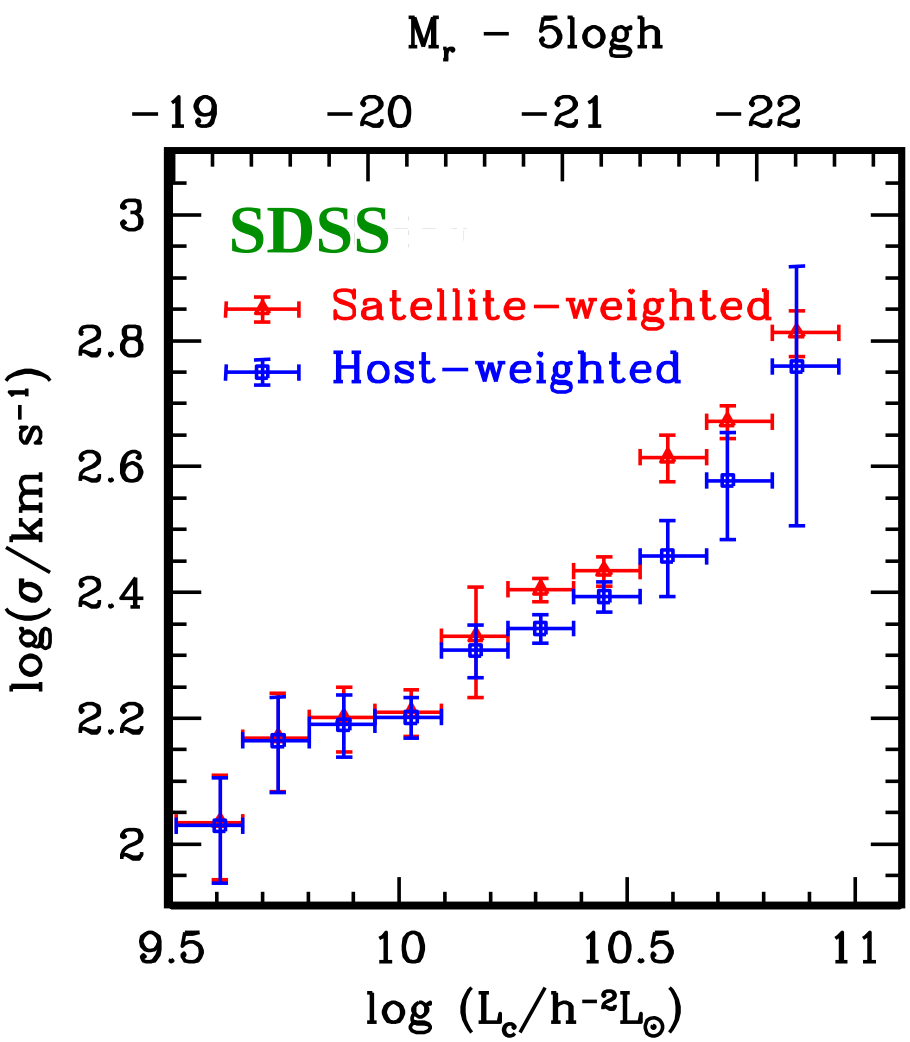
Satellite Weighting: each satellite receives weight of one

$$\sigma^2_{\rm sw} = \frac{\int P(M|L_c)\,\langle N_{\rm sat}\rangle_M\,\sigma^2_{\rm sat}(M)\,{\rm d}M}{\int P(M|L_c)\,\langle N_{\rm sat}\rangle_M\,{\rm d}M}$$

Host Weighting: each host receives weight of one

$$\sigma^2_{\rm hw} = \frac{\int P(M|L_c)\,\sigma^2_{\rm sat}(M)\,{\rm d}M}{\int P(M|L_c)\,{\rm d}M}$$

Satellite Kinematics in the SDSS



Based on SDSS volume-limited sample with 3863 centrals & 6101 satellites

Note that $\sigma_{\rm sw} \neq \sigma_{\rm hw} \Rightarrow$ non-zero scatter in $P(M|L_{\rm c})$

Modeling Methodology & Results

Recall:

$$\sigma_{\rm sw}^2 = \frac{\int P(M|L_c)\,\langle N_{\rm sat}\rangle_M\,\sigma_{\rm sat}^2(M)\,{\rm d}M}{\int P(M|L_c)\,\langle N_{\rm sat}\rangle_M\,{\rm d}M}$$

$$\sigma_{\rm hw}^2 = \frac{\int P(M|L_c)\,\sigma_{\rm sat}^2(M)\,{\rm d}M}{\int P(M|L_c)\,{\rm d}M}$$

- **Jeans equations yield $\sigma_{\rm sat}^2(M)$ for NFW halos**
- **$P(M|L_c)$ and $\langle N_{\rm sat}\rangle_M$ follow from CLF**
- **Constrain CLF model parameters by fitting the observed $\sigma_{\rm sw}(L_c)$ and $\sigma_{\rm hw}(L_c)$**

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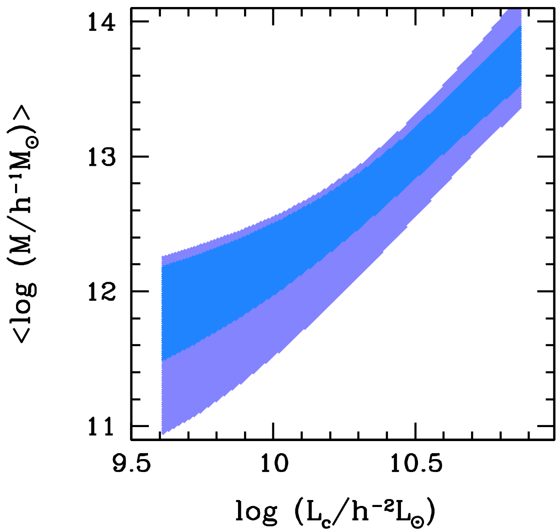

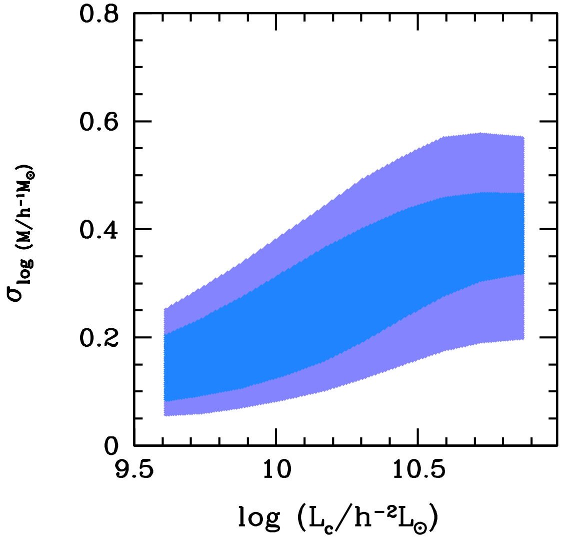

The 68 and 95 percent confidence levels from MCMC

Modeling Methodology & Results

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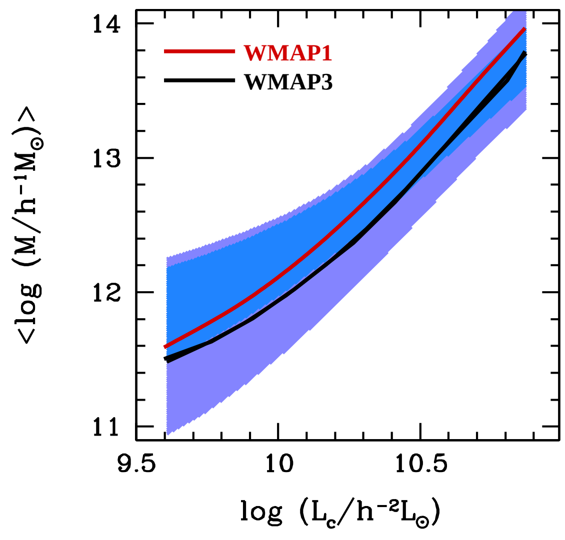
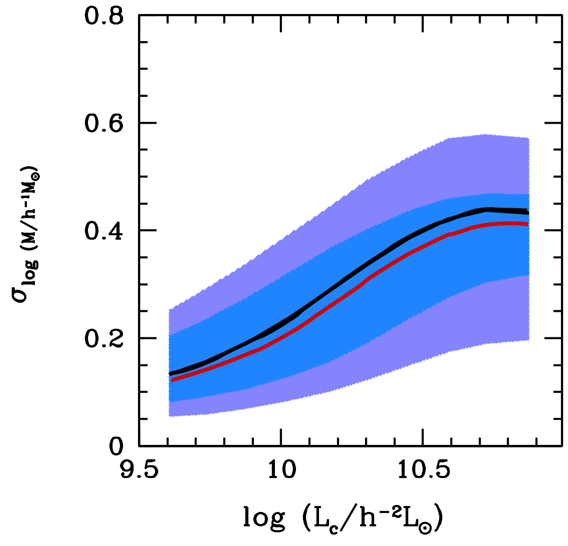
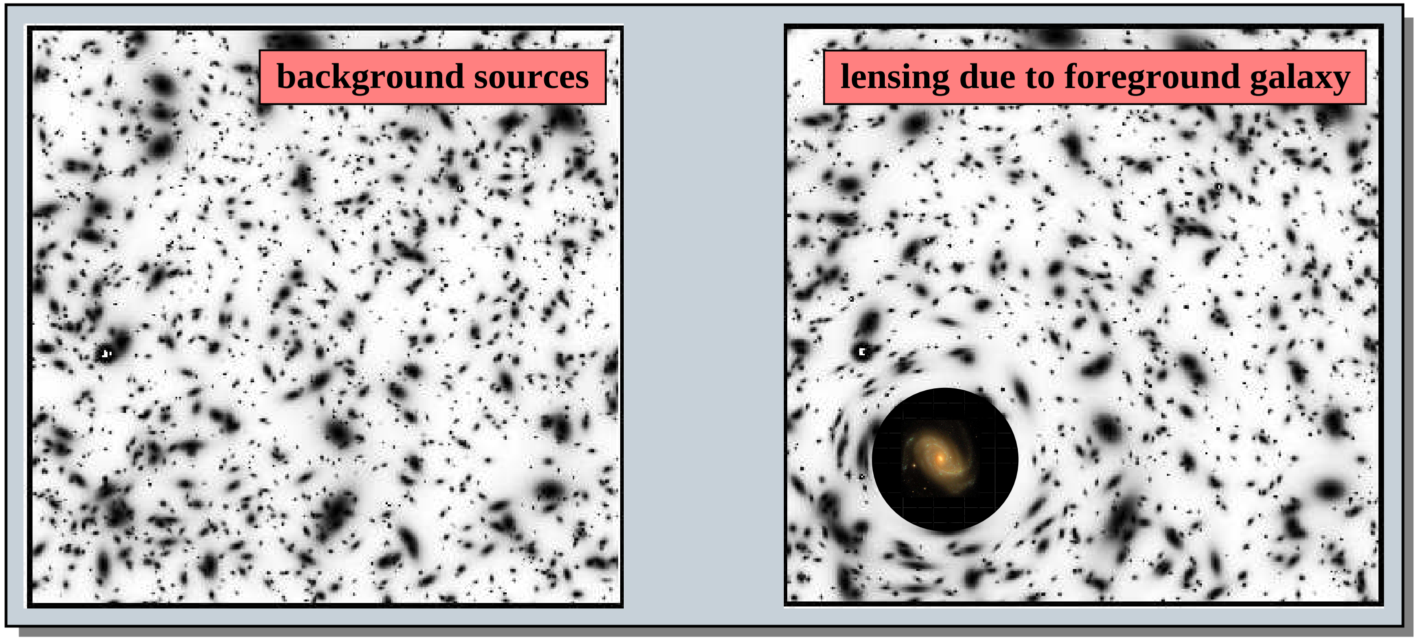




Good agreement with CLF clustering results

Galaxy-Galaxy Lensing

The mass associated with galaxies lenses background galaxies



Lensing causes correlated ellipticities, the tangential shear, γ_t, which is related to the excess surface density, $\Delta\Sigma$, according to

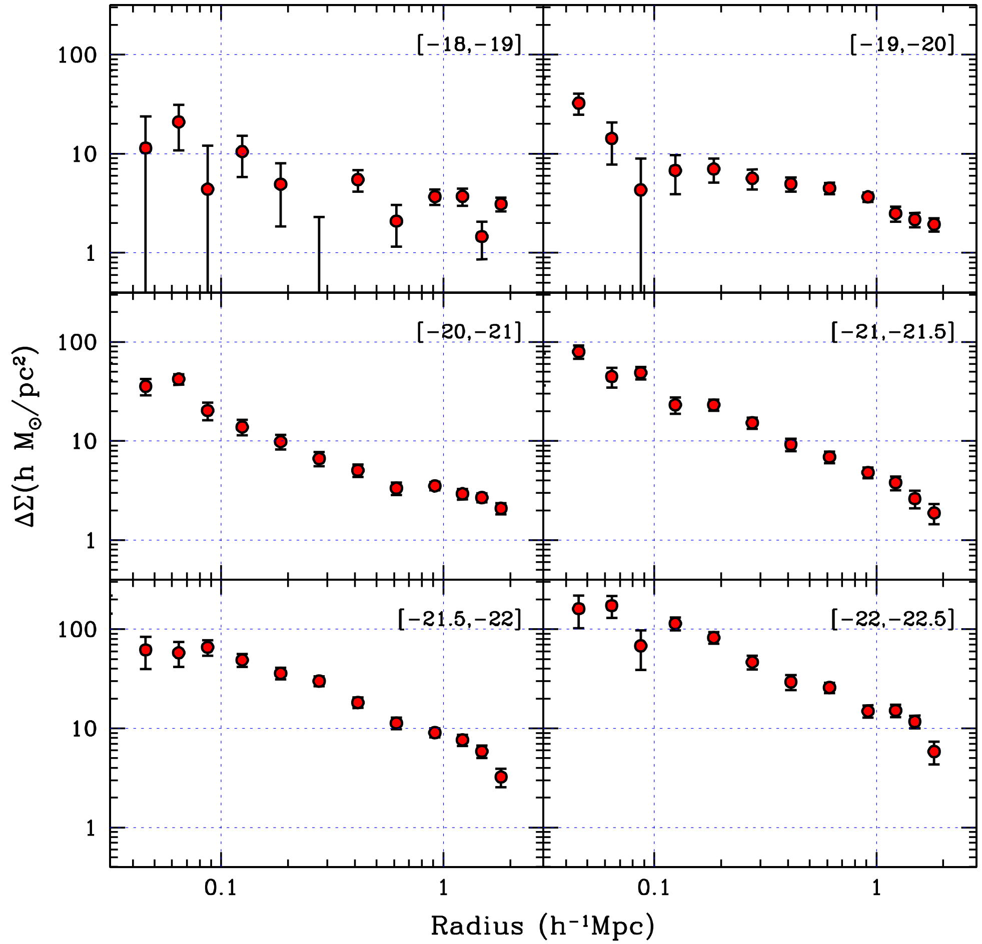
$$\gamma_t(R)\Sigma_{\rm crit} = \Delta\Sigma(R) = \bar{\Sigma}(<R) - \Sigma(R)$$

$\Sigma(R)$ is line-of-sight projection of galaxy-matter cross correlation:

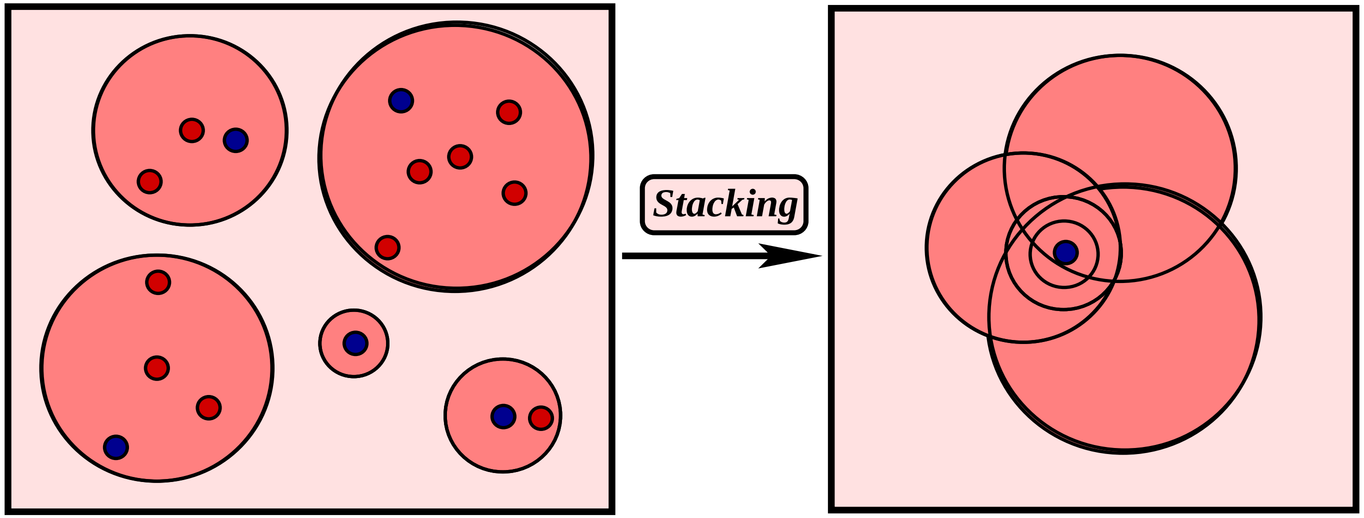
$$\Sigma(R) = \bar{\rho}\int_0^{D_S}\left[1+\xi_{\rm g,dm}(r)\right]d\chi$$

The Measurements

- **Number of background sources per lens is limited.**
- **Measuring γ_t with sufficient S/N requires stacking of many lenses**
- **$\Delta\Sigma(R|L_1, L_2)$ has been measured using the SDSS by Mandelbaum et al. (2005) for different bins in lens luminosity**



How to interpret the signal?



Because of stacking the lensing signal is difficult to interpret

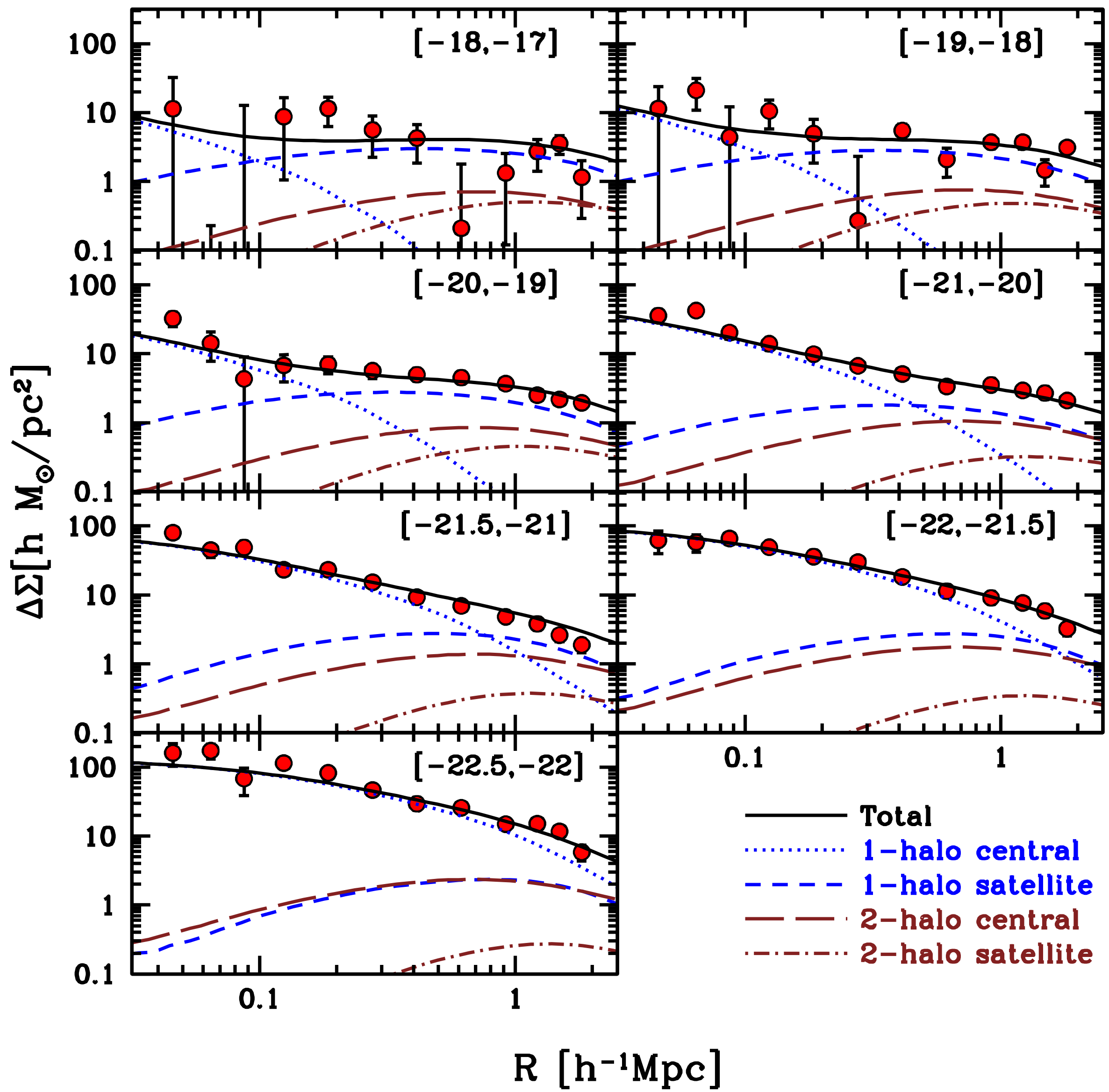
In order to model the data, what is required is:

$$P_{\rm cen}(M|L) \qquad P_{\rm sat}(M|L) \qquad f_{\rm sat}(L)$$

These can all be computed from the CLF

Using $\Phi(L|M)$ constrained from clustering data, we can predict the lensing signal $\Delta\Sigma(R|L_1, L_2)$

Comparison with CLF Predictions



NOTE: This is not a fit, but a prediction based on CLF

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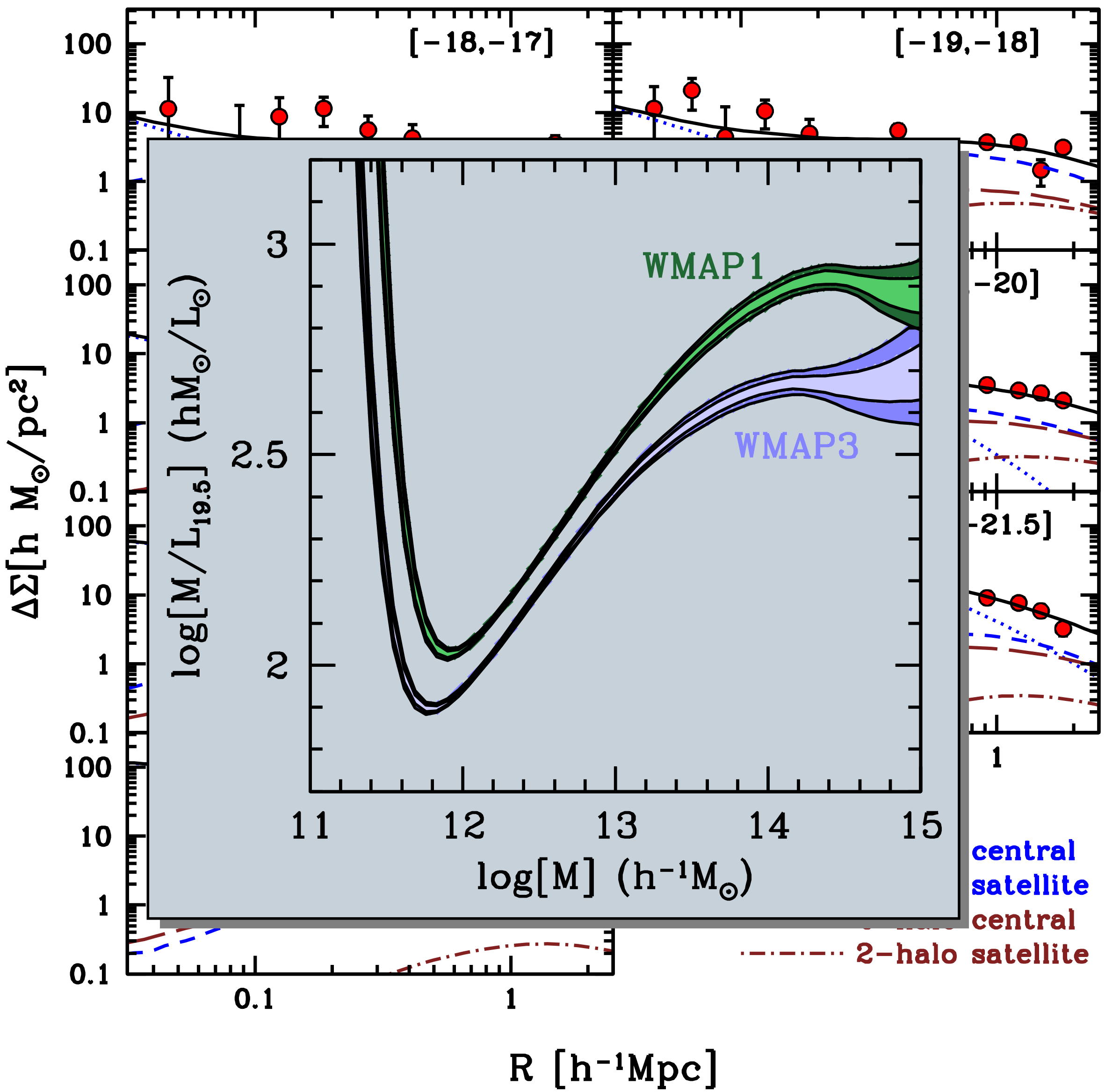

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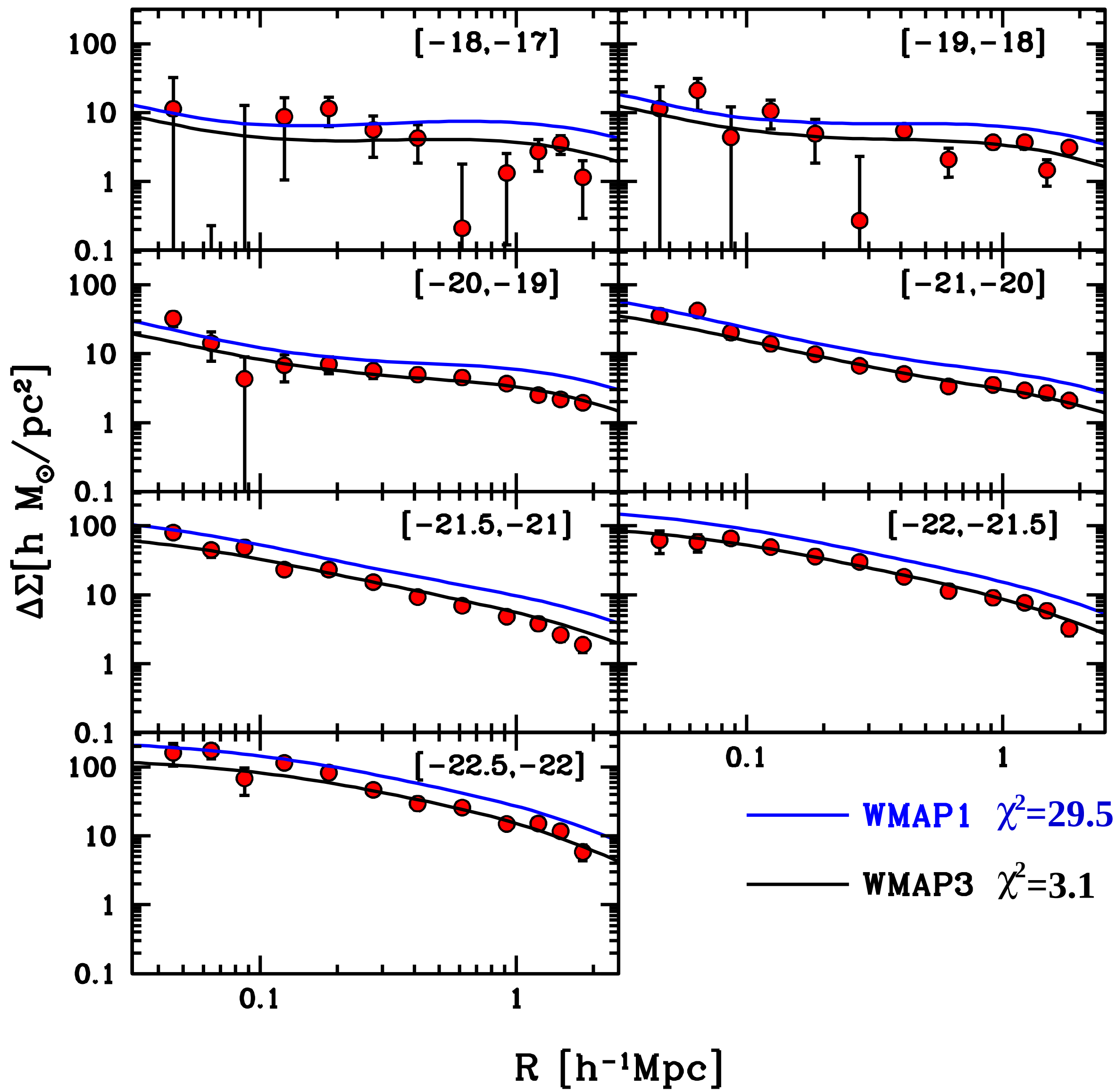
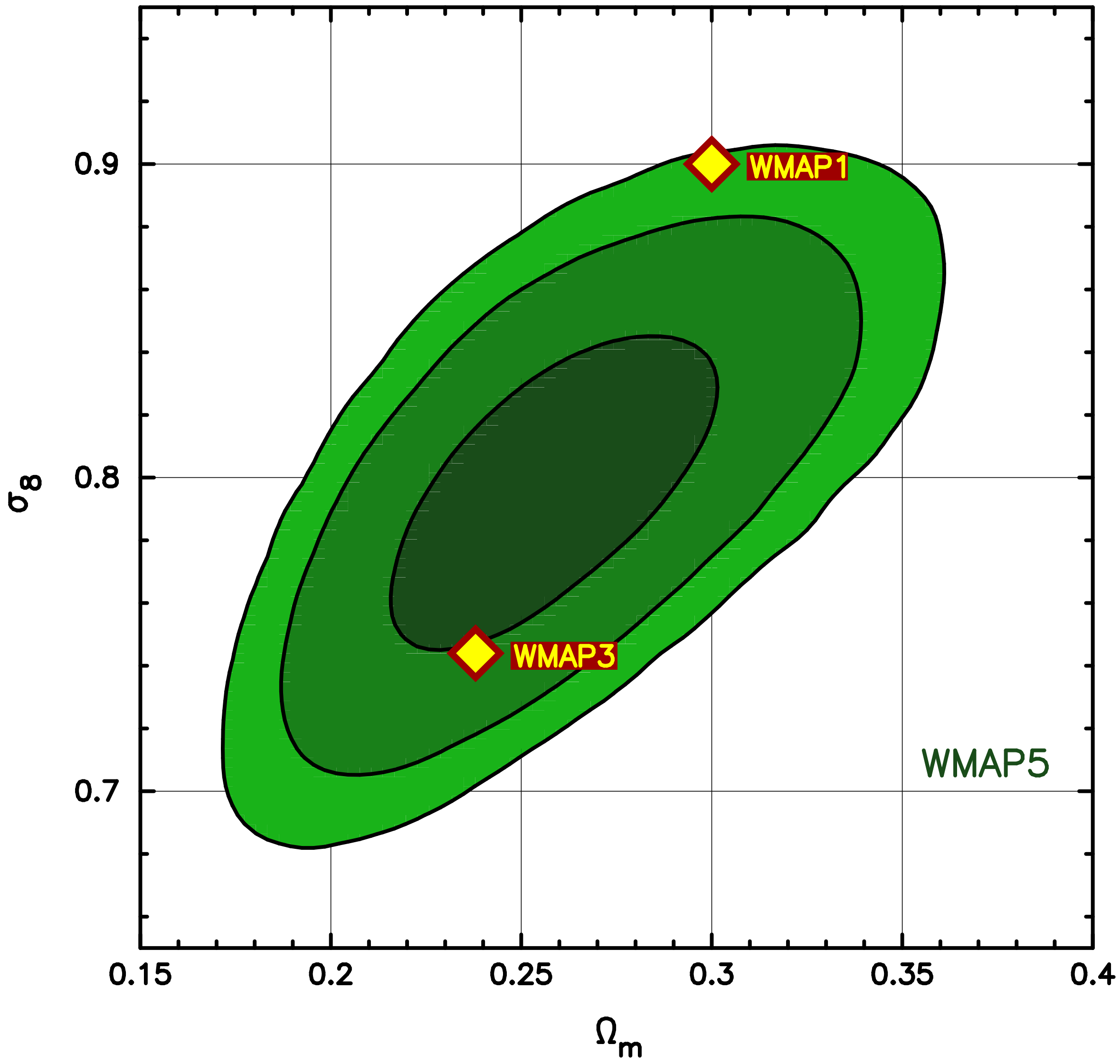
WMAP3 vs. WMAP1

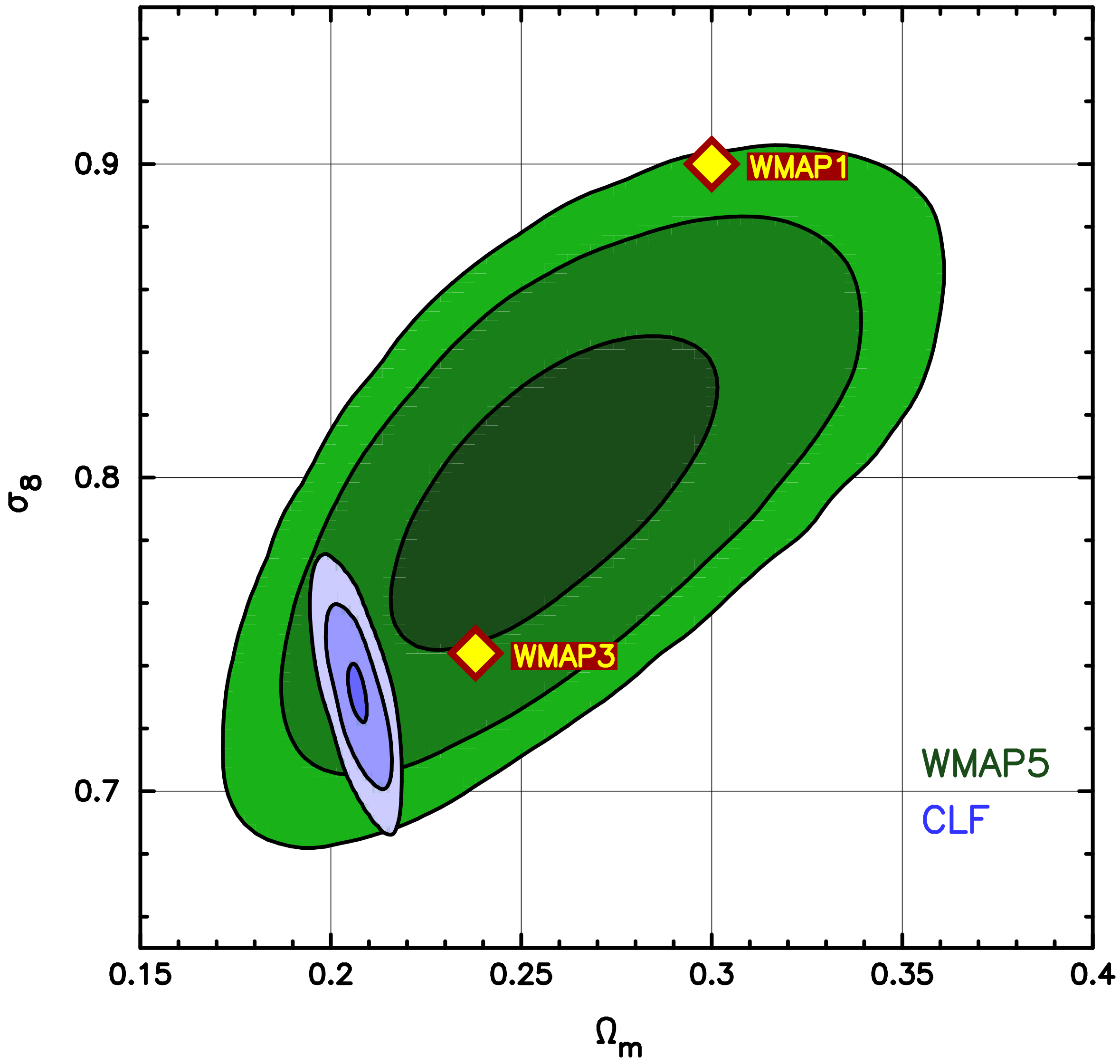


WMAP3 cosmology clearly preferred over WMAP1 cosmology

Cosmological Constraints



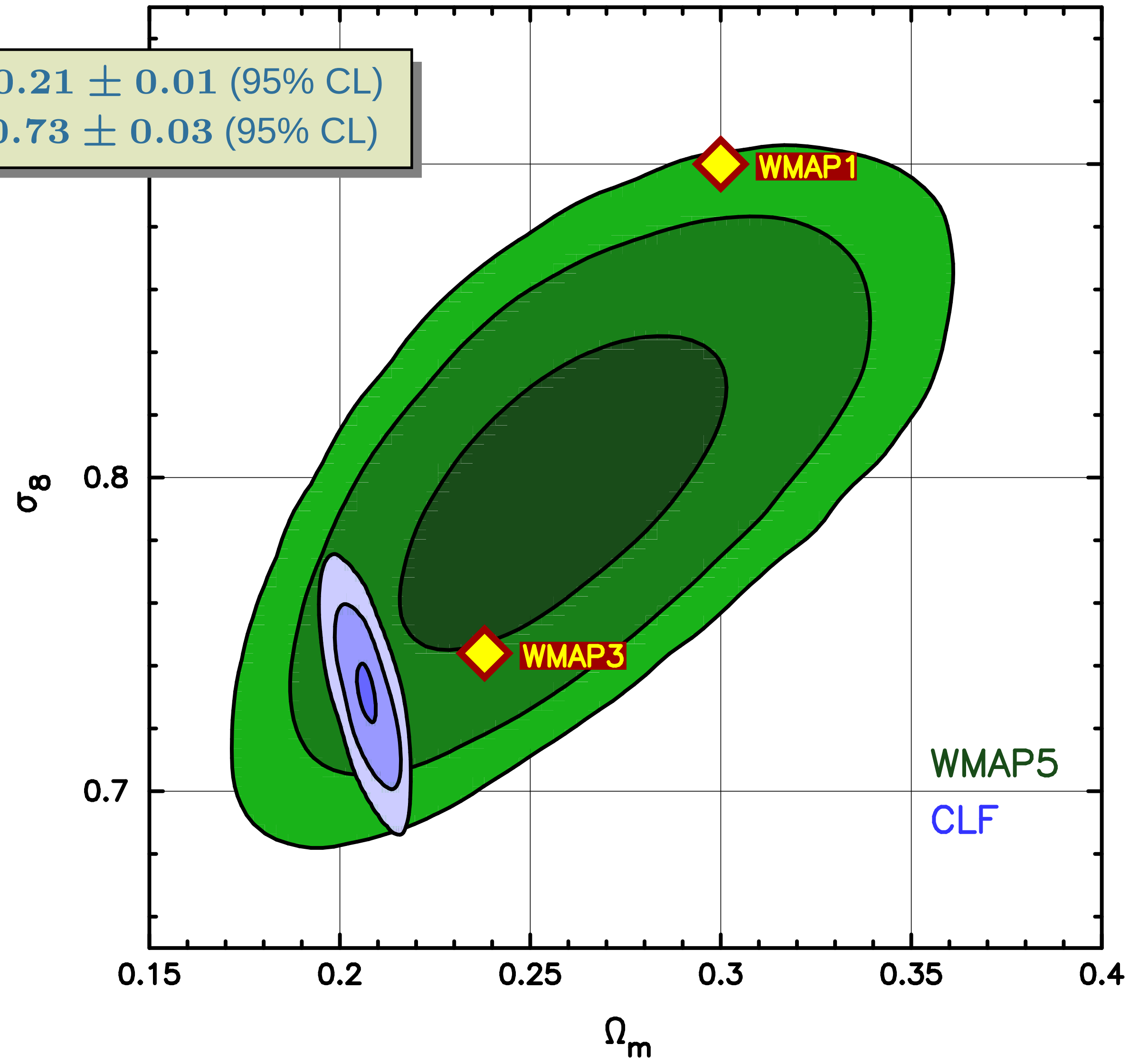
Cosmological Constraints



Precision Cosmology using non-linear structure!!

Cosmological Constraints

$\Omega_m = 0.21 \pm 0.01$ (95% CL)
$\sigma_8 = 0.73 \pm 0.03$ (95% CL)



Precision Cosmology using non-linear structure!!

Conclusions

Four methods to statistically constrain $P(M|L)$

Group Catalogues	**Clustering**
Satellite Kinematics	**Galaxy-Galaxy Lensing**

- **Requires somewhat arbitrary group-finder**
- **We used well-tested Halo Based Group Finder of Yang et al (2005)**
- **Ideal for studying environment dependence of galaxy formation**
- **Mass assignments is cosmology-dependent**
- **Correlation function of groups is direct reflection of that of dark matter haloes**

Conclusions

Four methods to statistically constrain $P(M|L)$

Group Catalogues	**Clustering**
Satellite Kinematics	**Galaxy-Galaxy Lensing**

- **Straightforward to constrain $P(M|L)$ with CLF**
- **Accurate constraints from large galaxy redshift surveys**
- **Results are strongly cosmology-dependent**

Conclusions

Four methods to statistically constrain $P(M|L)$

Group Catalogues	**Clustering**
Satellite Kinematics	**Galaxy-Galaxy Lensing**

- **Requires selection of centrals and satellites from redshift surveys**
- **Requires stacking and is therefore sensitive to scatter in $P(M|L)$**
- **Using satellite weighting and host weighting simultaneously constrains both mean and scatter of $P(M|L)$**
- **Even with large redshift surveys such as SDSS, statistics are limited**
- **Data not sufficient to discriminate between WMAP1 and WMAP3**

Conclusions

Four methods to statistically constrain $P(M|L)$

Group Catalogues	Clustering
Satellite Kinematics	Galaxy-Galaxy Lensing

- Lensing probes masses directly
- Requires stacking and is therefore sensitive to scatter in $P(M|L)$
- Also very sensitive to satellite fractions $f_{\rm sat}(L)$
- Most easily interpreted with use of CLF $\Phi(L|M)$
- Combination of lensing and clustering holds potential to tightly constrain cosmological parameters

Cosmological Conclusions

Cosmological constraints obtained from non-linear structure (clustering + lensing + group catalogue) are in excellent agreement with CMB constraints

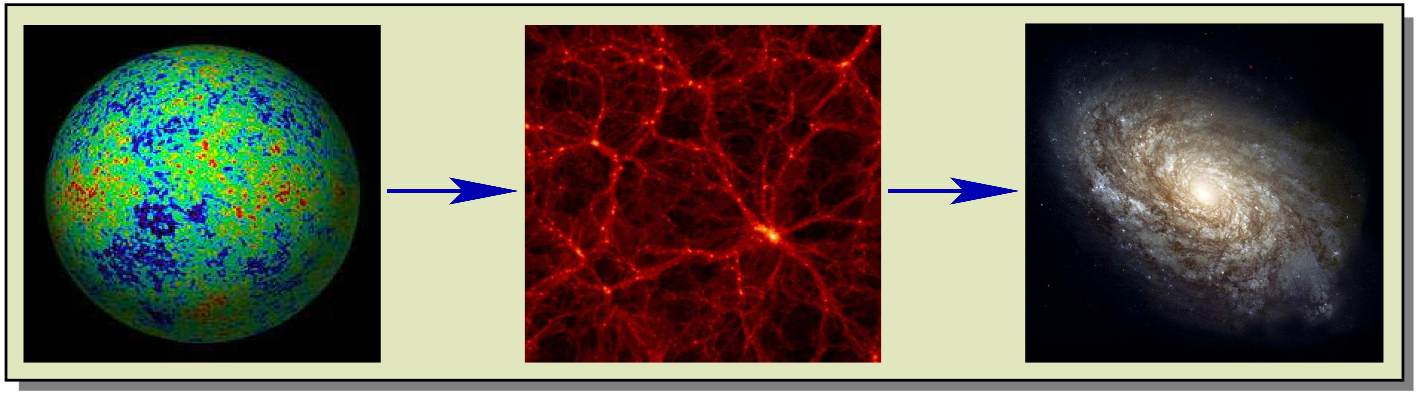
Current (preliminary) results suggest

$$\Omega_m = 0.21 \pm 0.01 \text{ (95\% CL)}$$
$$\sigma_8 = 0.73 \pm 0.03 \text{ (95\% CL)}$$

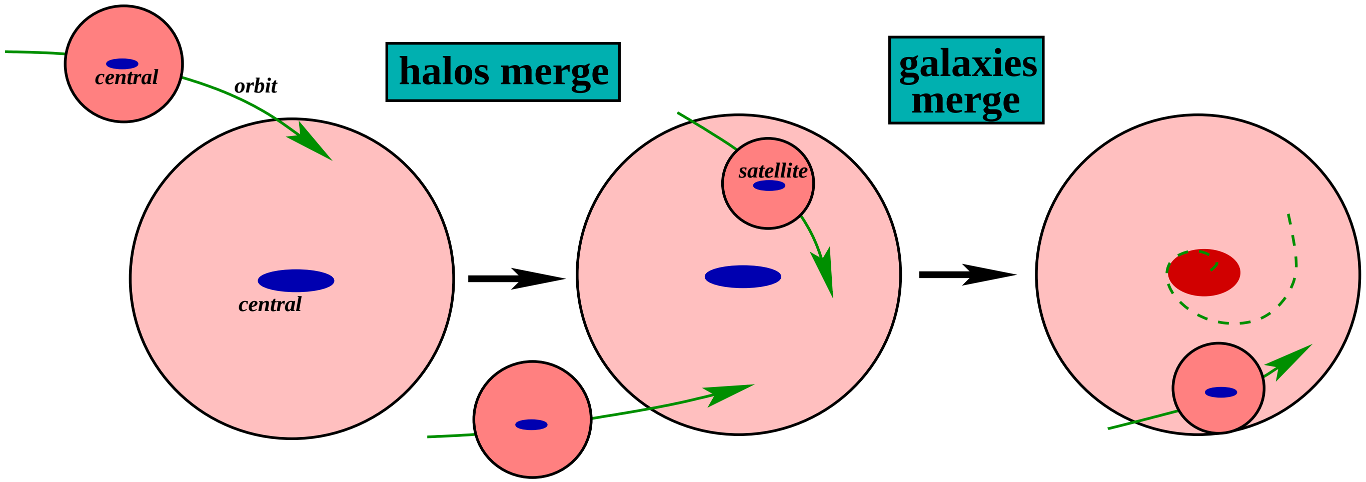
This technique is competative with and complementary to BAO, cosmic shear, SNIa and Lyα forest

If anything, our results indicate that our model for structure formation is accurate on non-linear scales

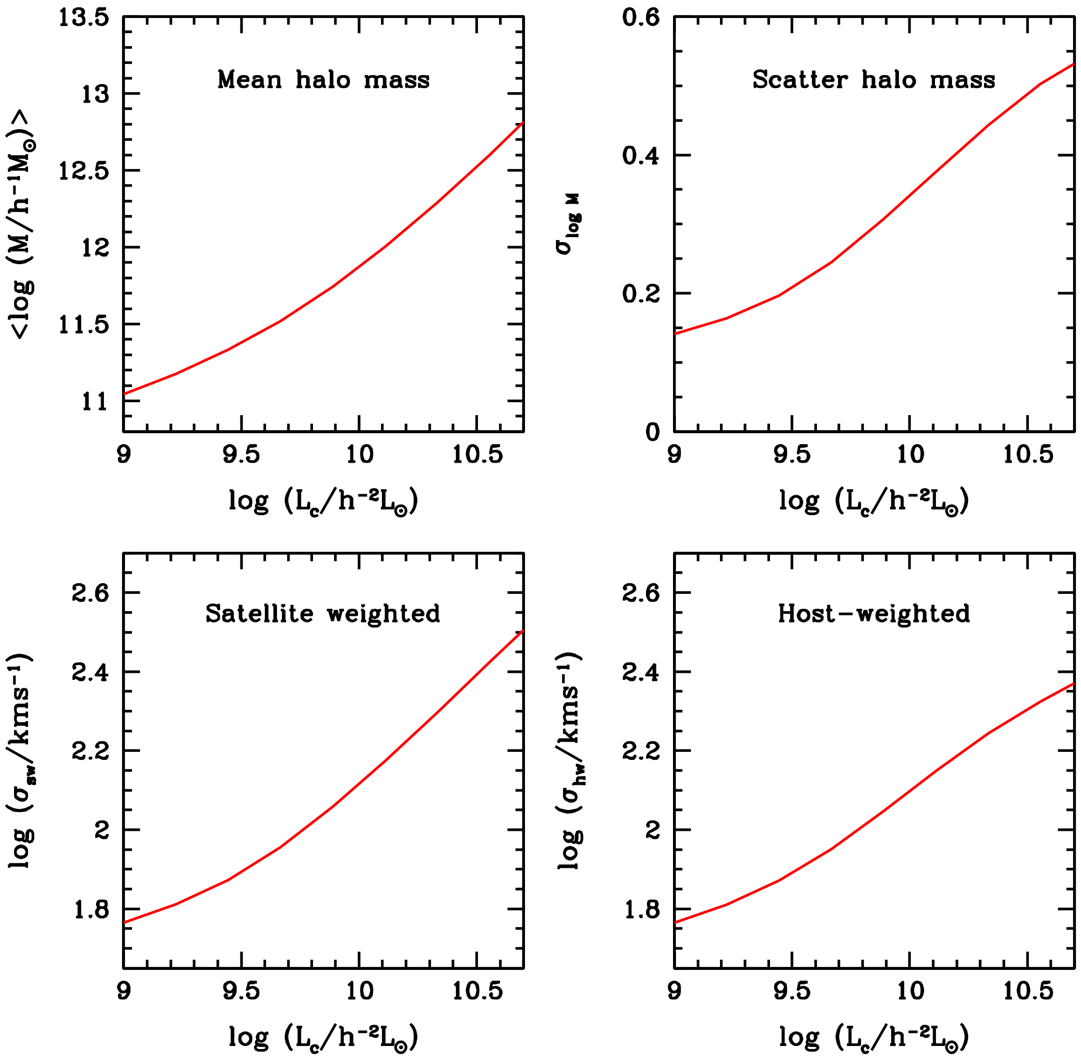
Galaxy Formation in a Nutshell



- **Perturbations grow due to gravitational instability and collapse to produce (virialized) dark matter halos**
- **Baryons cool, accumulate at center, and form stars ⇒ galaxy**
- **Dark matter halos merge, causing hierarchical growth**
- **Halo mergers create satellite galaxies that orbit halo**



Satellite Weighting or Host Weighting?



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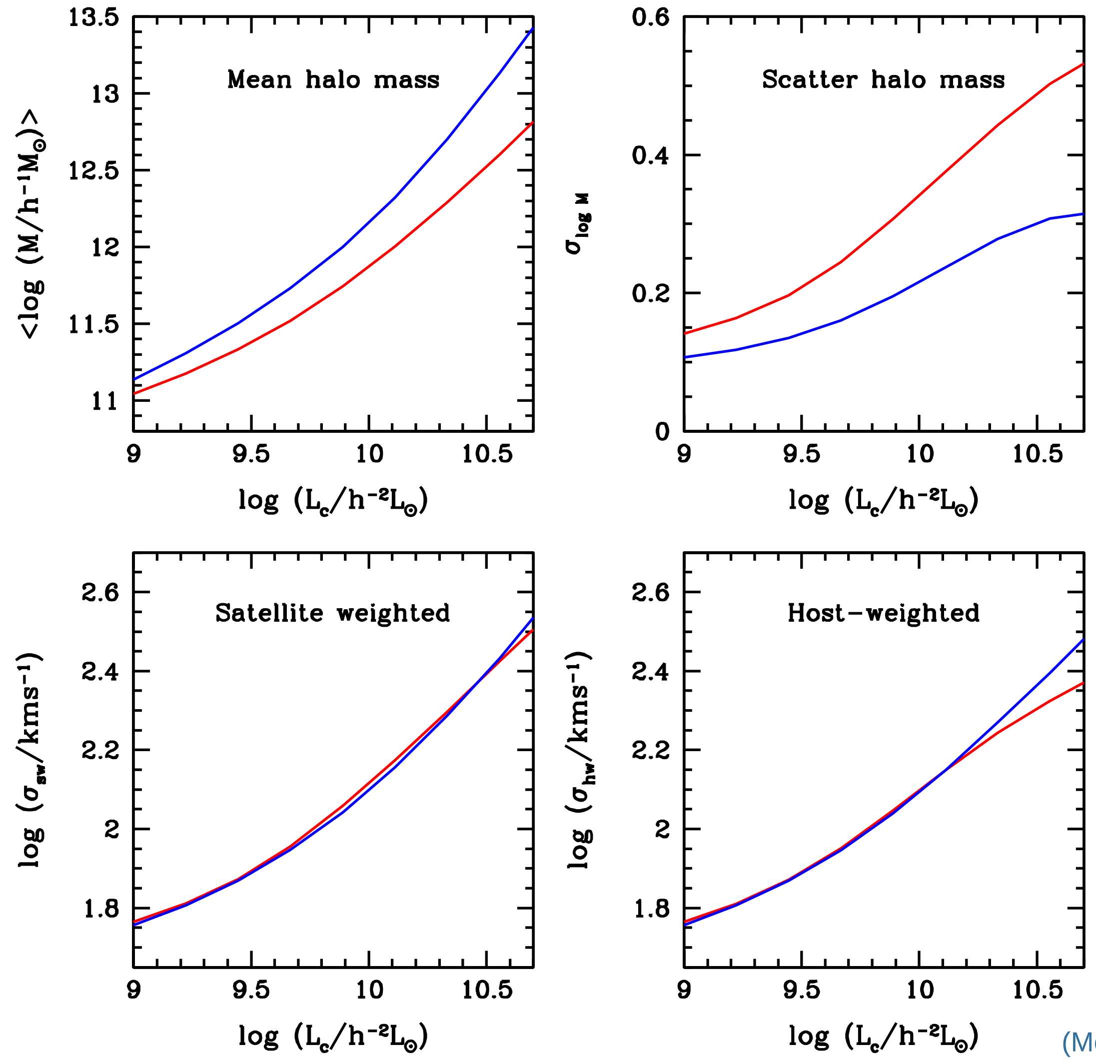
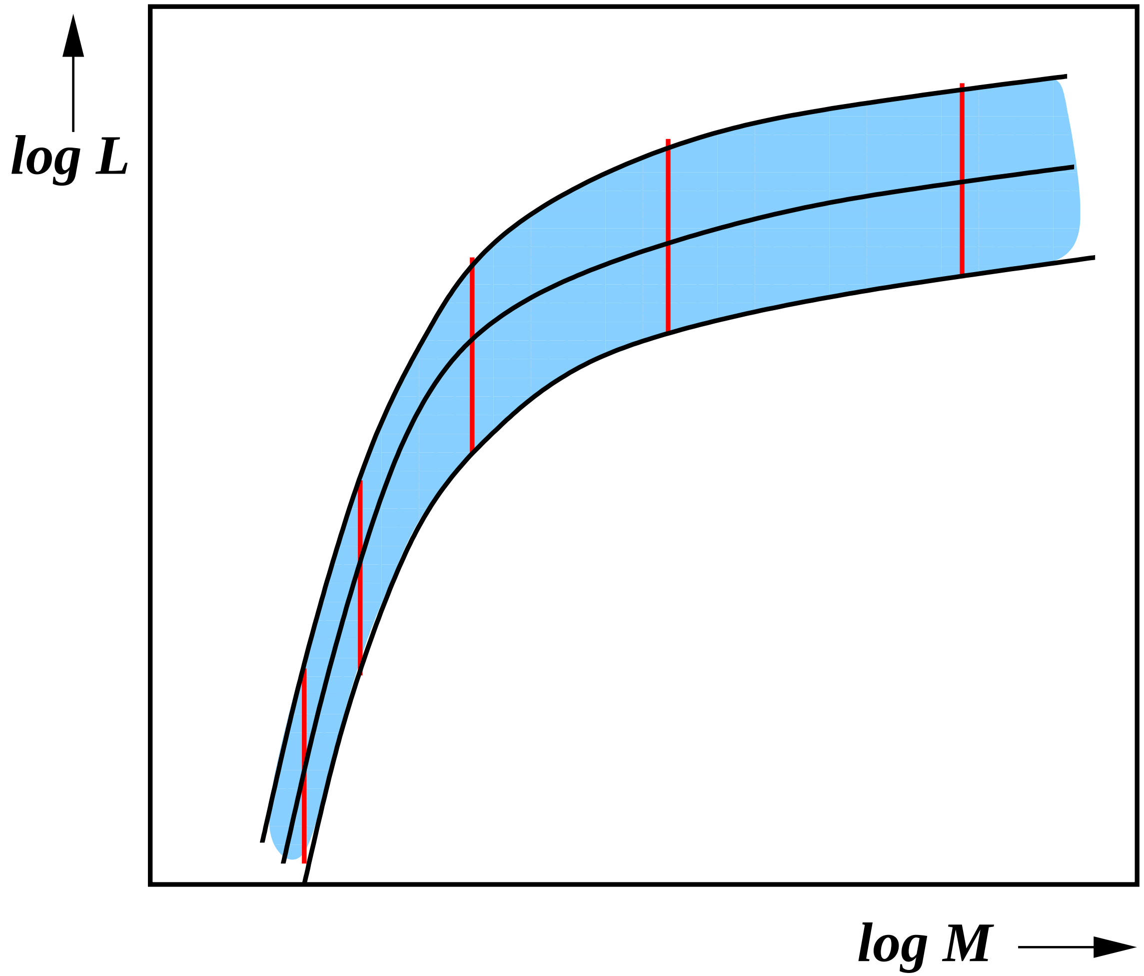


(More, vdB et al. 2008)

The combination of $\sigma_{\rm sw}$ and $\sigma_{\rm sw}$ allows one to determine mean and scatter of $P(M|L_{\rm c})$

Implications for Galaxy Formation Stochasticity



- The scatter in $P(L_{\rm cen}|M)$ is independent of M

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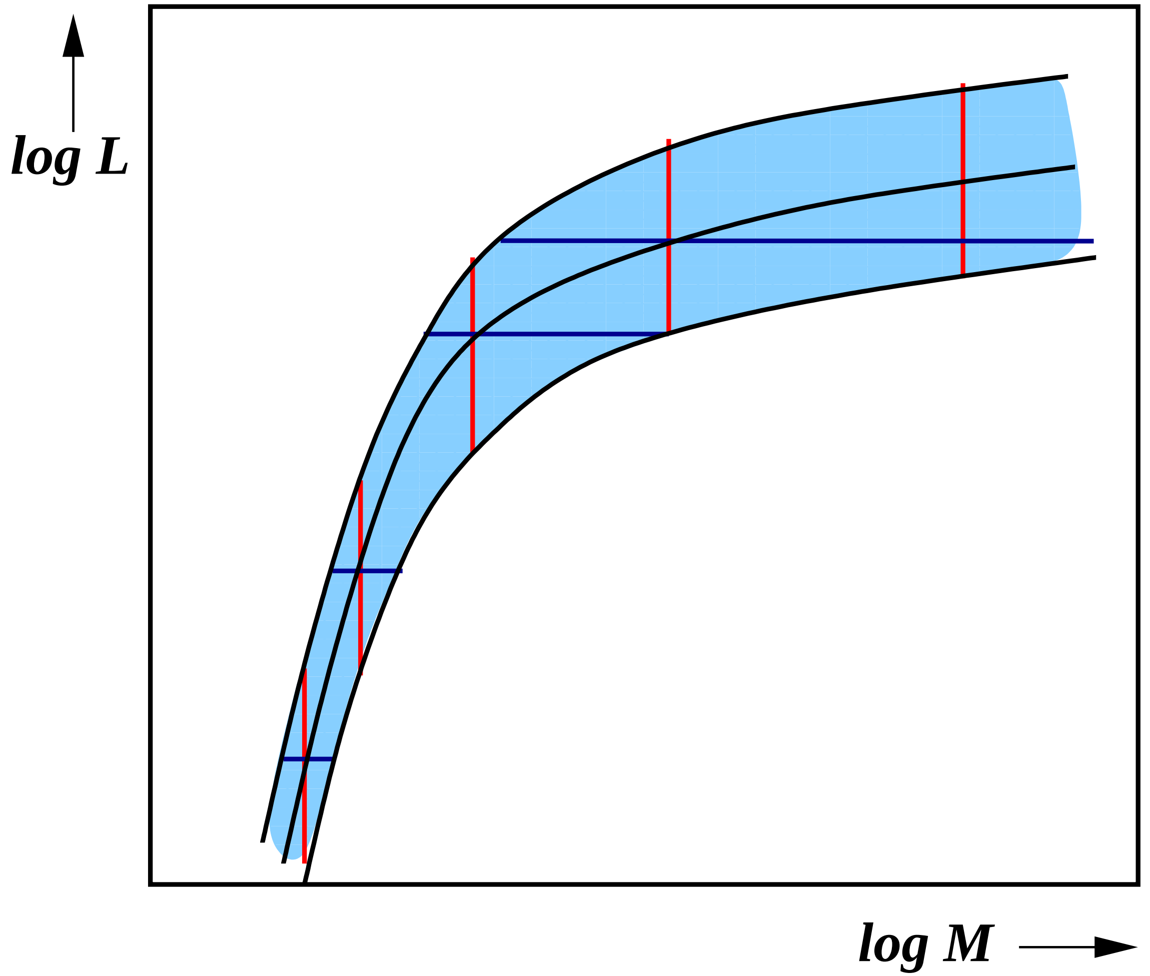
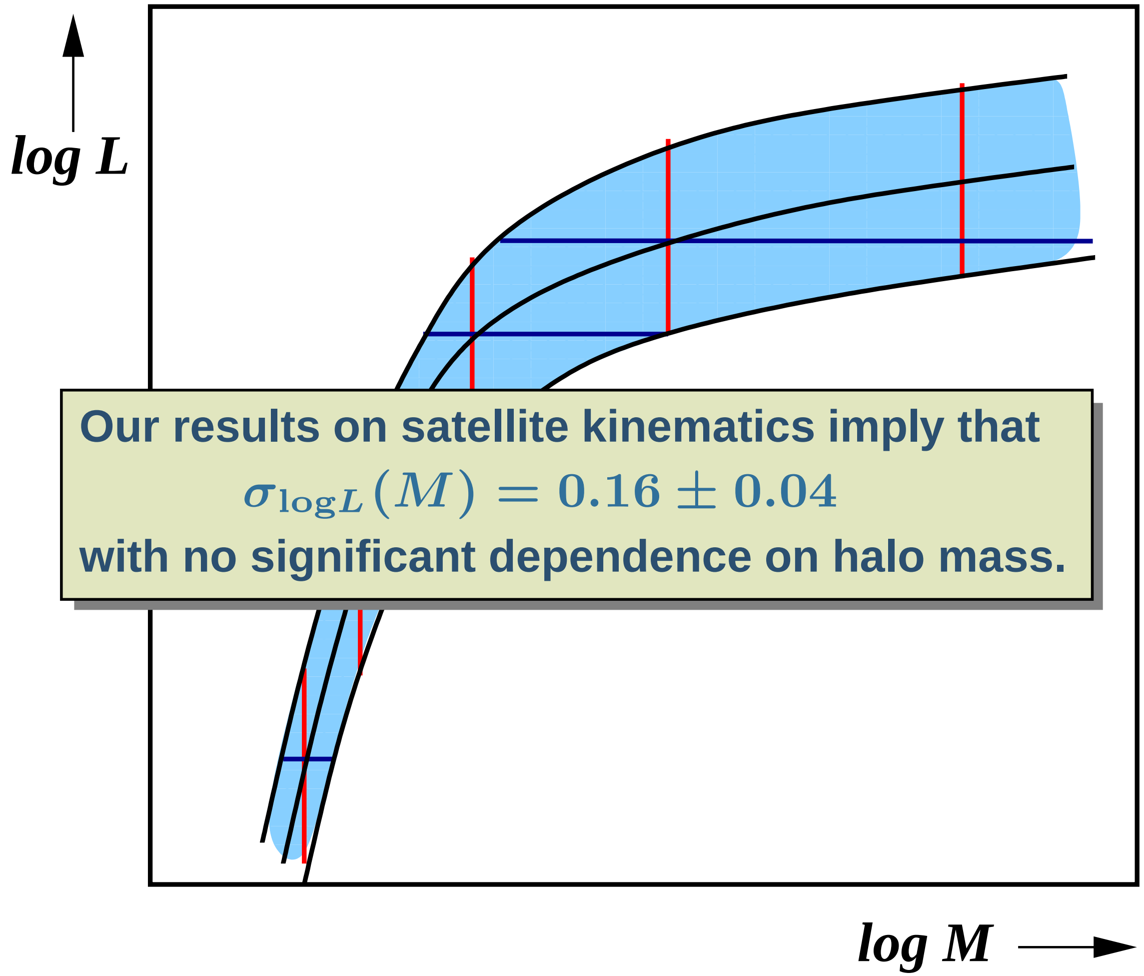


- The scatter in $P(L_{\rm cen}|M)$ is independent of M
- The scatter in $P(M|L_{\rm cen})$ increases strongly with $L_{\rm cen}$

Implications for Galaxy Formation Stochasticity



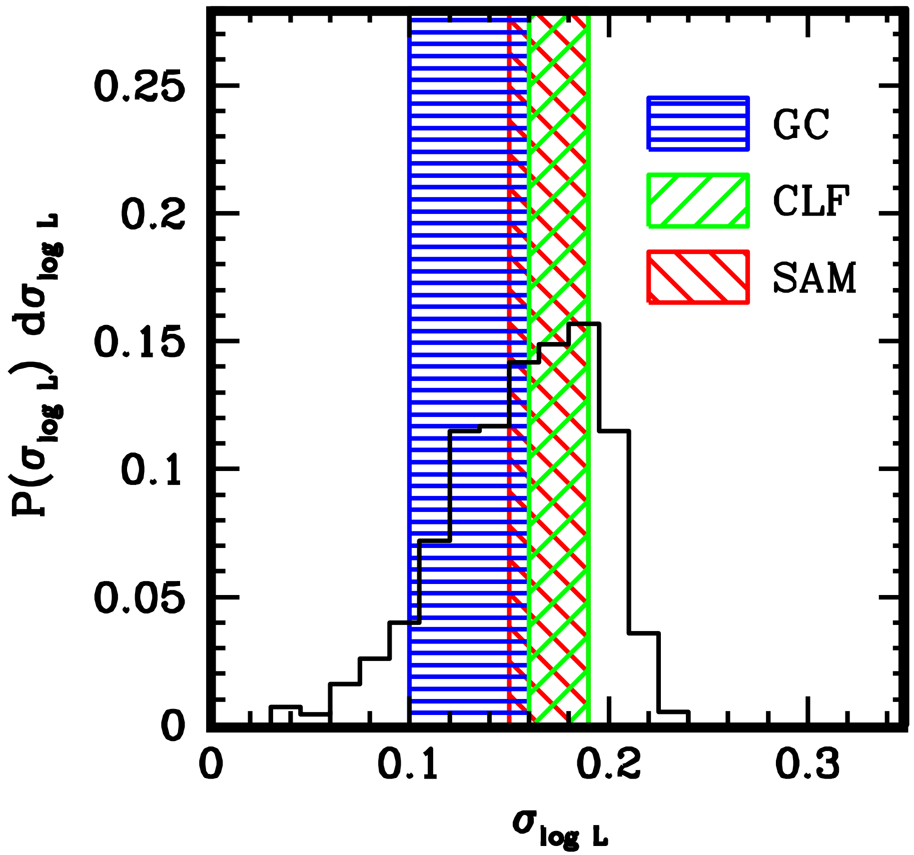
Our results on satellite kinematics imply that

$$\sigma_{\log L}(M) = 0.16 \pm 0.04$$

with no significant dependence on halo mass.

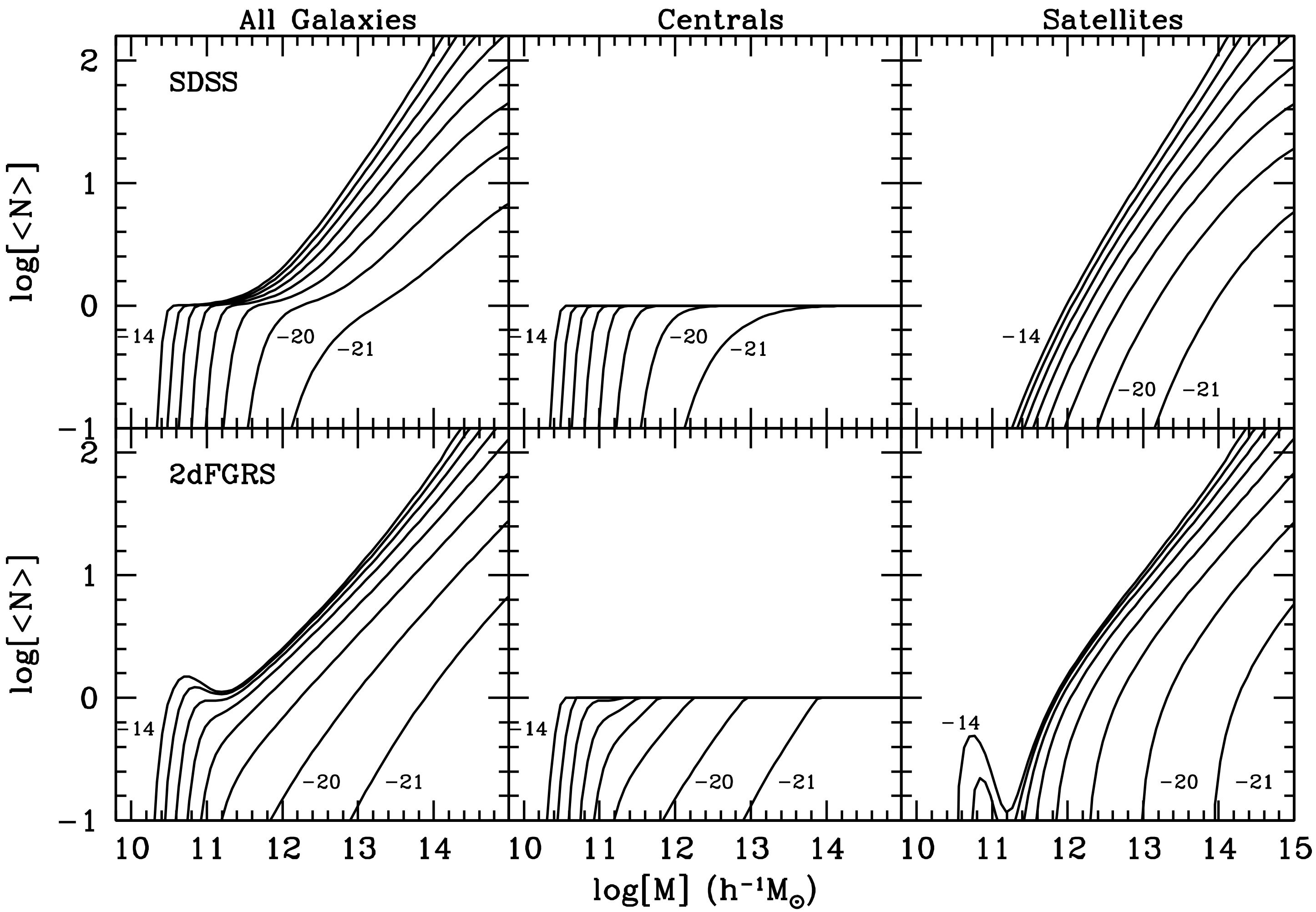
- **The scatter in $P(L_{\rm cen}|M)$ is independent of M**
- **The scatter in $P(M|L_{\rm cen})$ increases strongly with $L_{\rm cen}$**

Comparison with other Constraints



- **Probability Distribution from Satellite Kinematics**
- **Constraints from Galaxy Group Catalogue (Yang et al. 2008)**
- **Constraints from Clustering Analysis (Cooray 2006)**
- **Predictions from Semi Analytical Model (Croton et al. 2006)**

Halo Occupation Numbers



- **Unlike 2dFGRS, the SDSS reveals clear shoulders at $\langle N \rangle_M = 1$**
- **Most likely this is an ‘artefact’ of the functional form of the CLF**