### The Role of Magnetic Fields in Population III Star Formation

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#### "What about magnetic fields?"

#### Increasingly well-explored space ...

#### Xu, et al (2008)

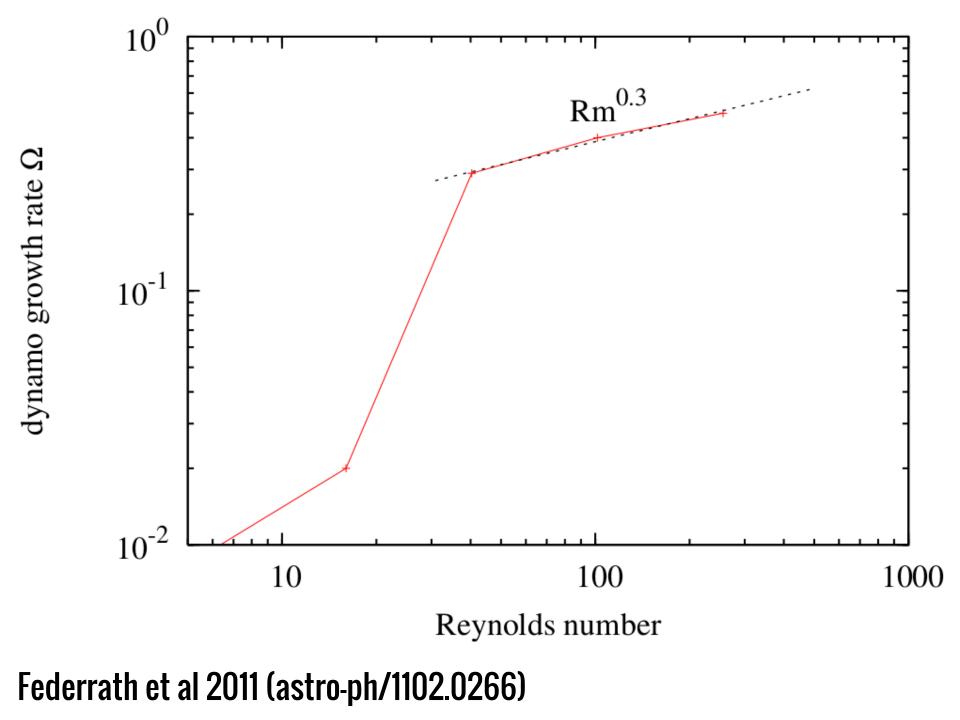
Biermann battery: <u>not</u> sufficient for dynamical importance!

(astro-ph/0807.2647)

#### Federrath, et al (2011)

## Small-scale dynamo action on scales less than Jeans Length

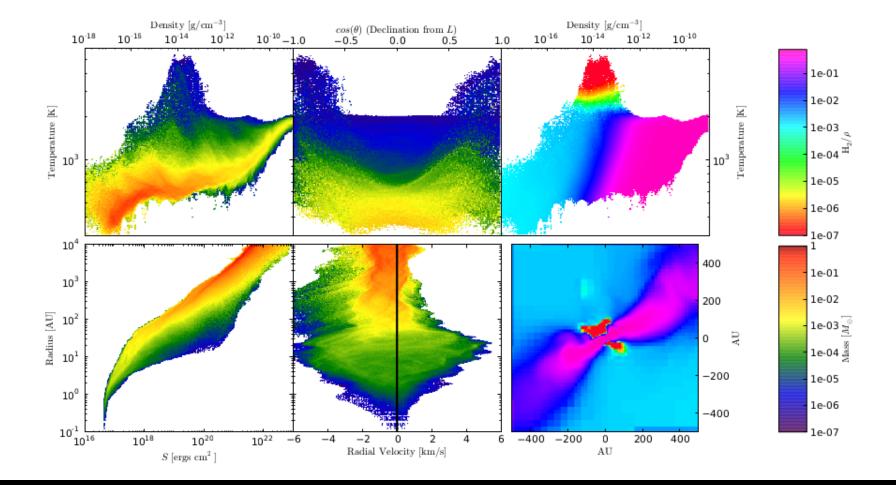
(astro-ph/1102.0266)



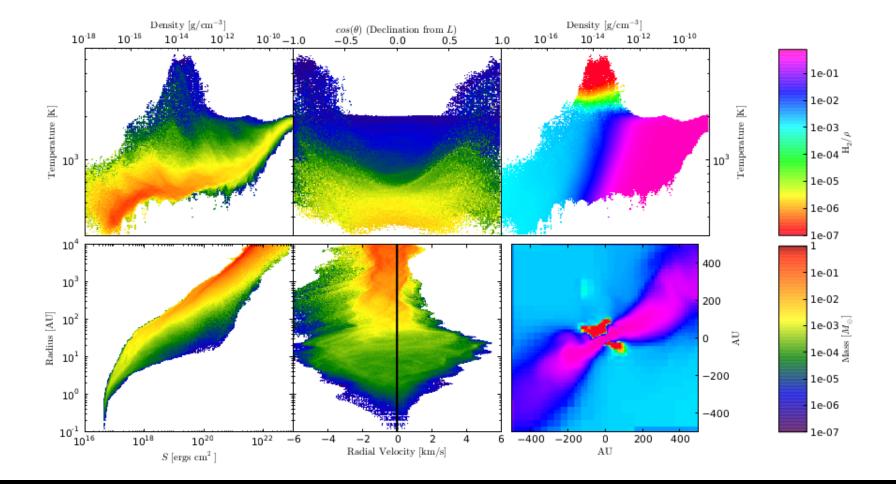
#### Schober, et al (2012)

#### Saturation of magnetic field energy from small-scale dynamo

(astro-ph/1204.0658)



#### Do magnetic fields alter disk structure?

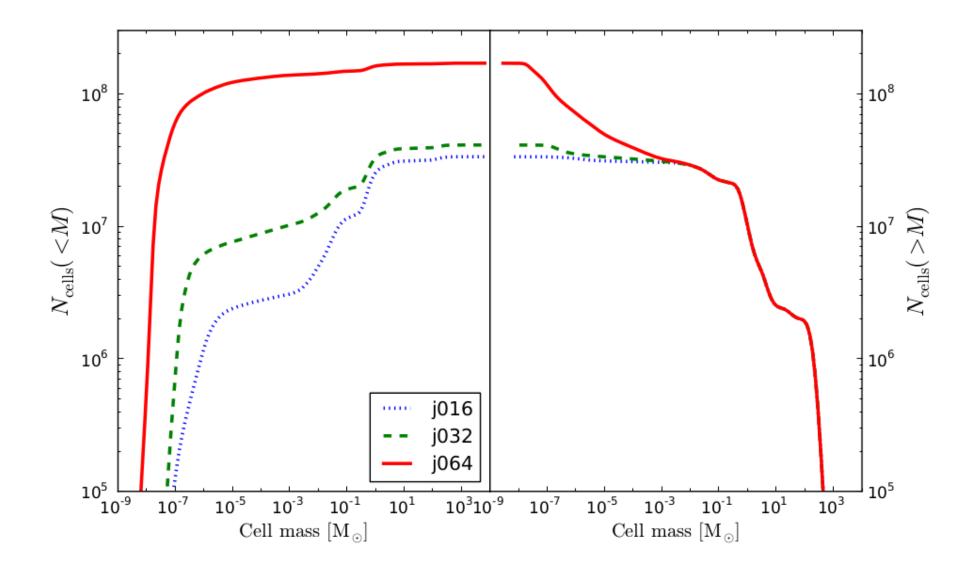


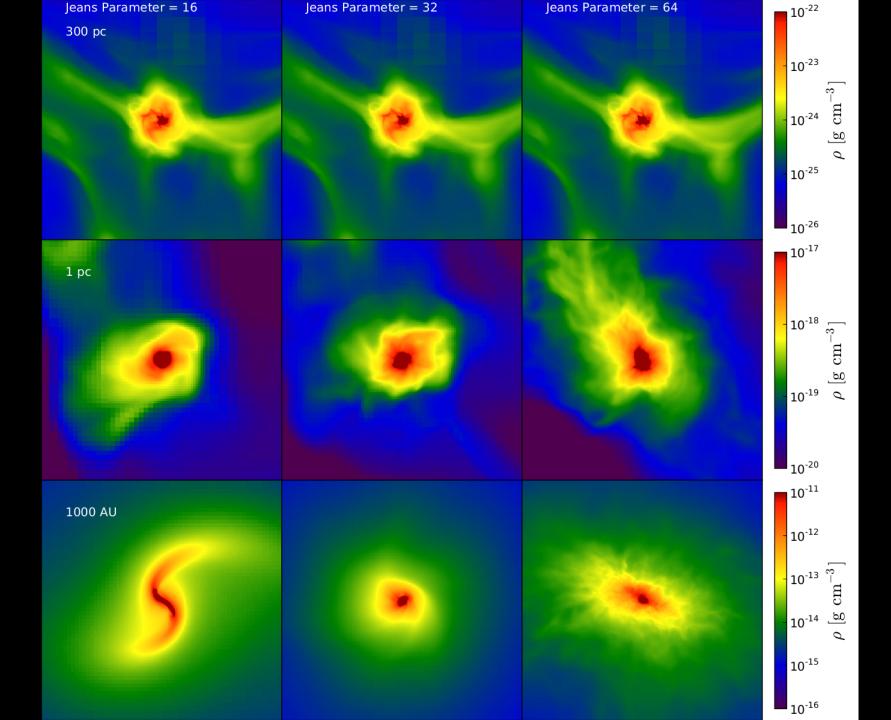
#### Do we see amplification in cosmology?

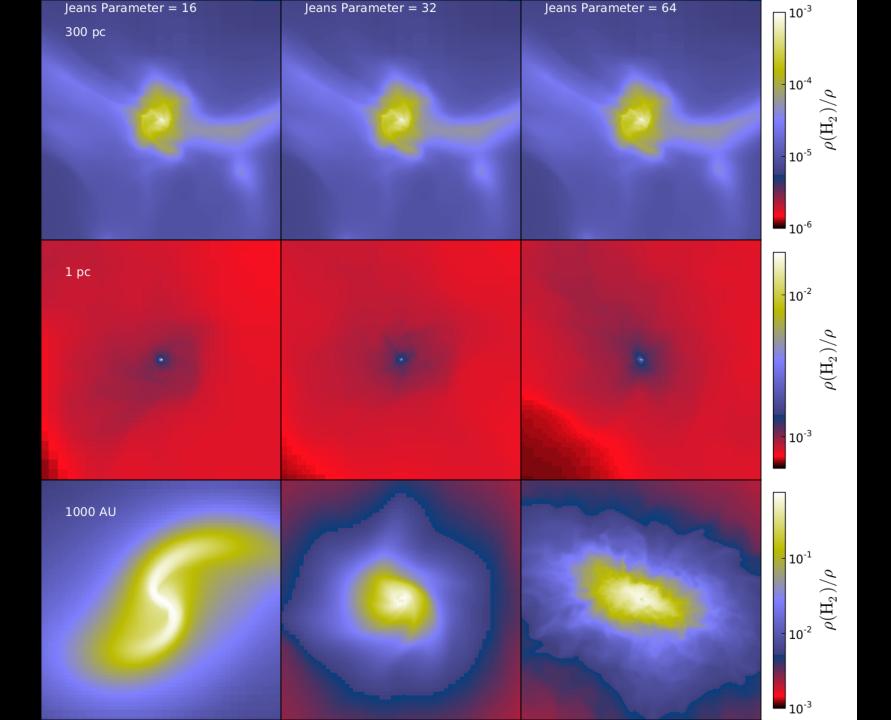
#### Seed field of $10^{-14}$ G at z = 99.

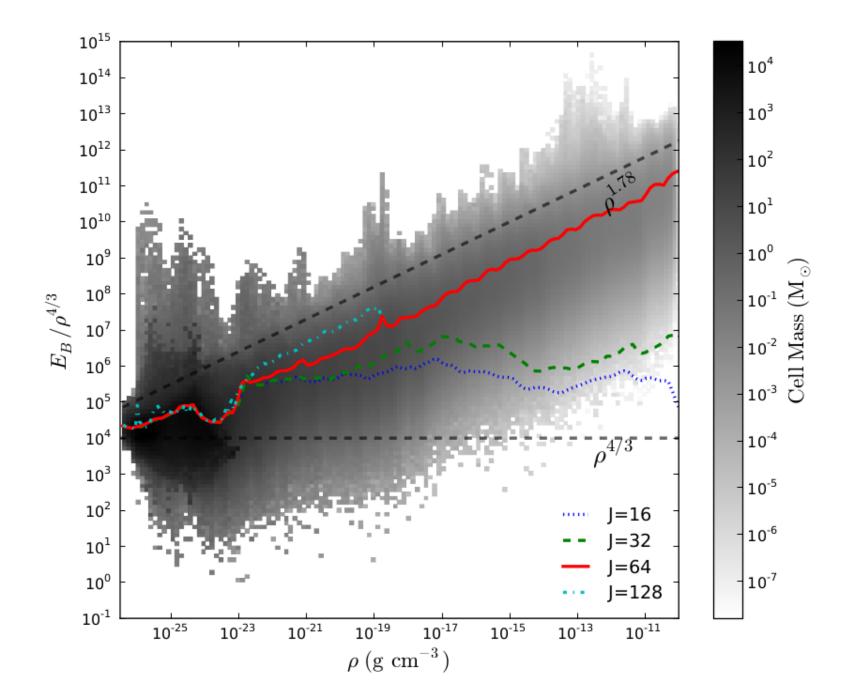
Resolve the 200 K Jeans Length by 16, 32, 64 cells.

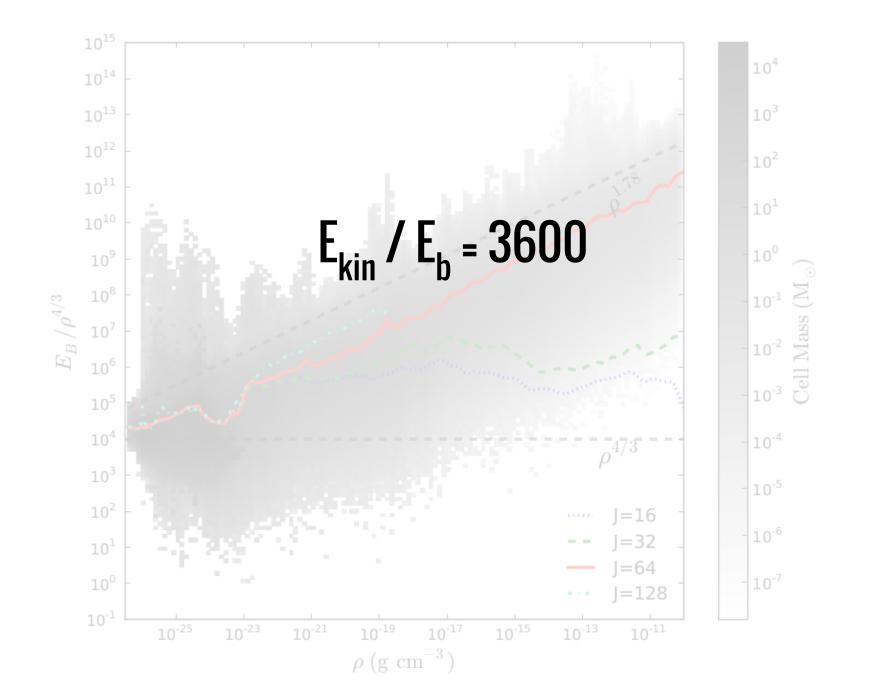
(astro-ph/1112.4479)

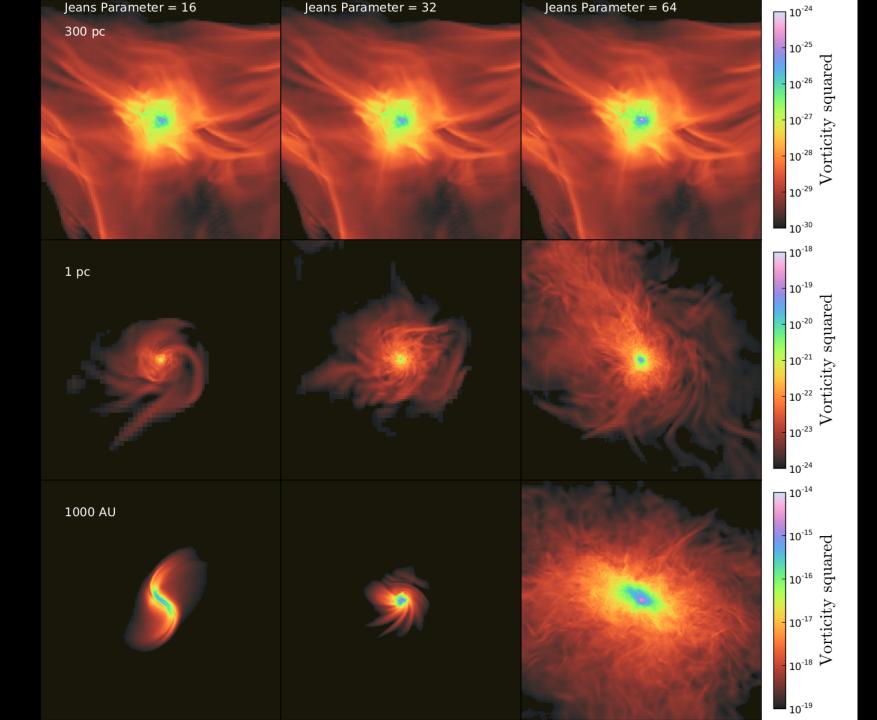


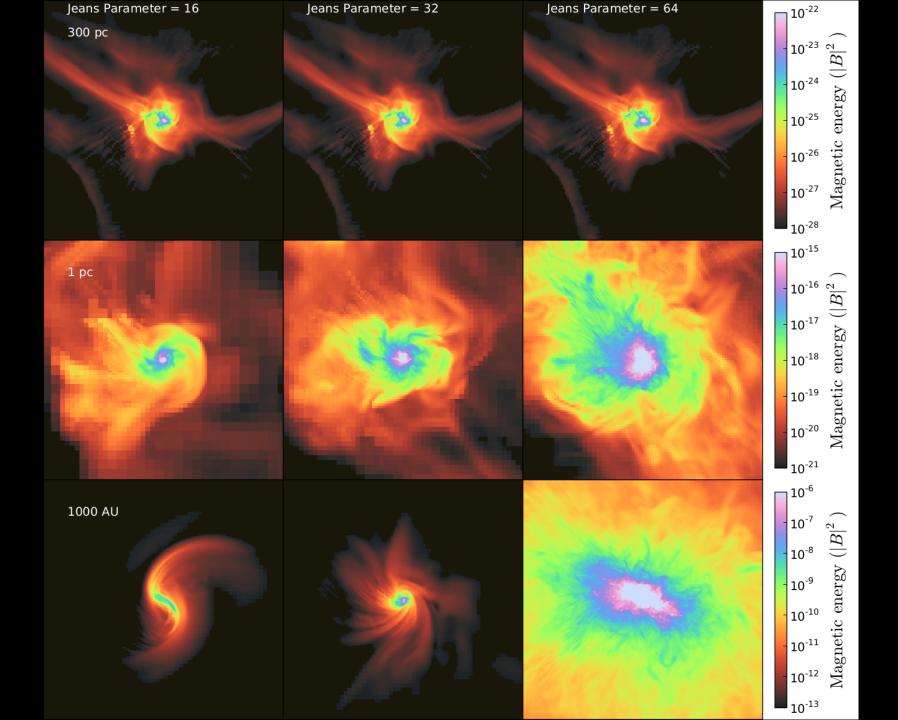


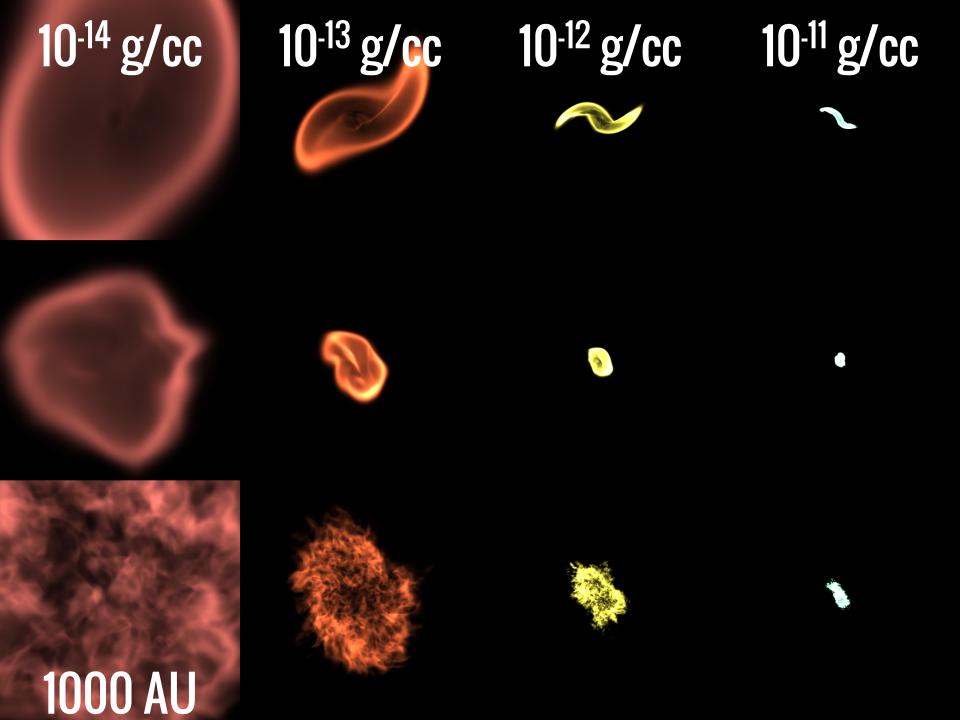
















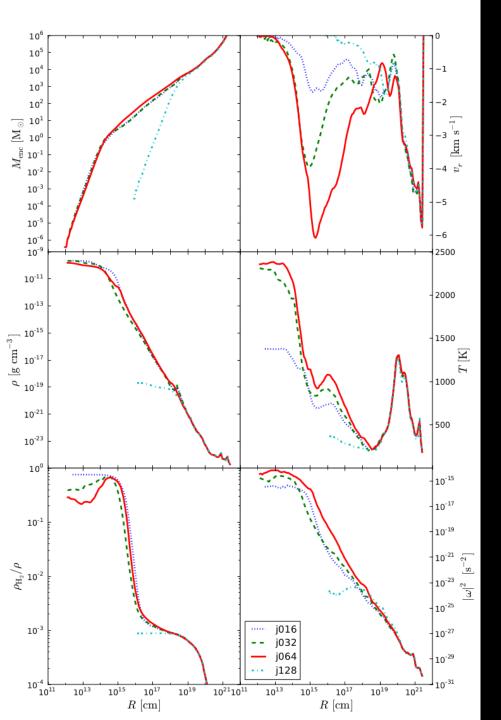




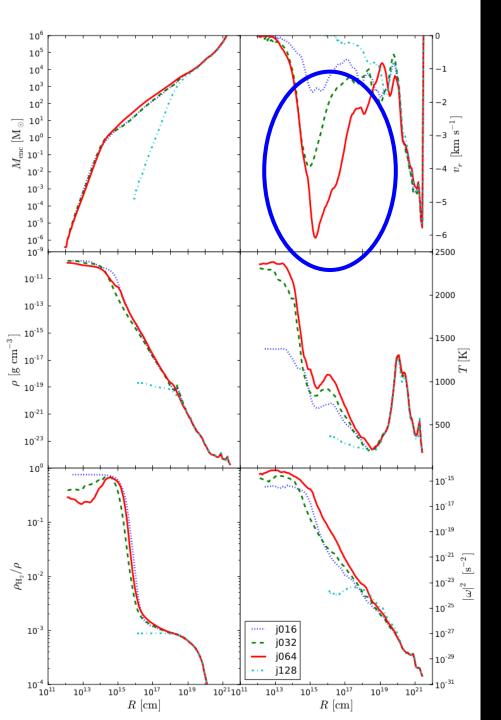




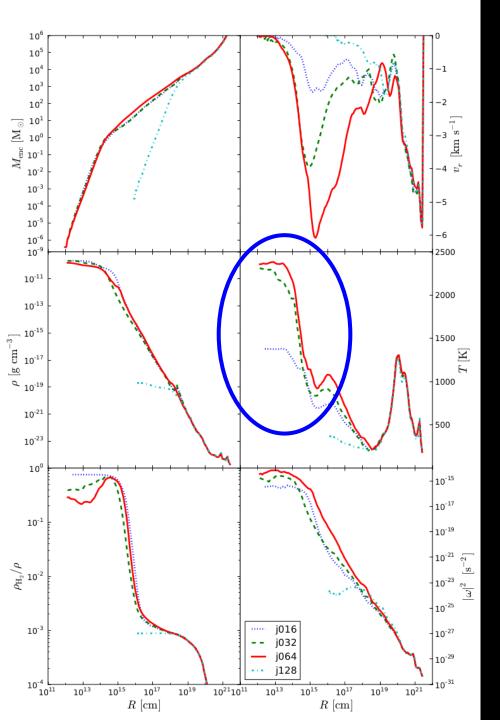
### 1000 AU



#### Three primary

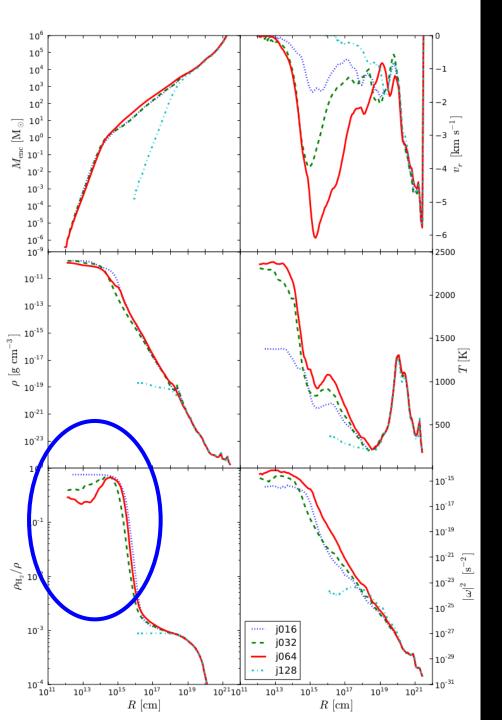


#### **Increased inward velocity**



#### Increased inward velocity

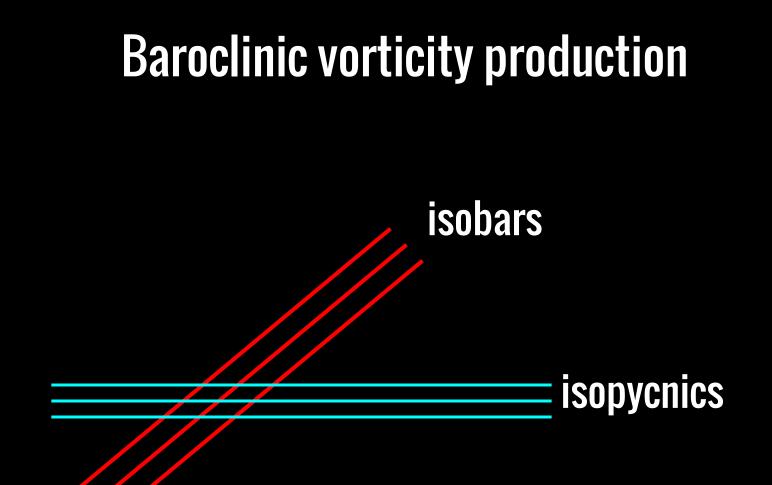
#### Increased core temperature



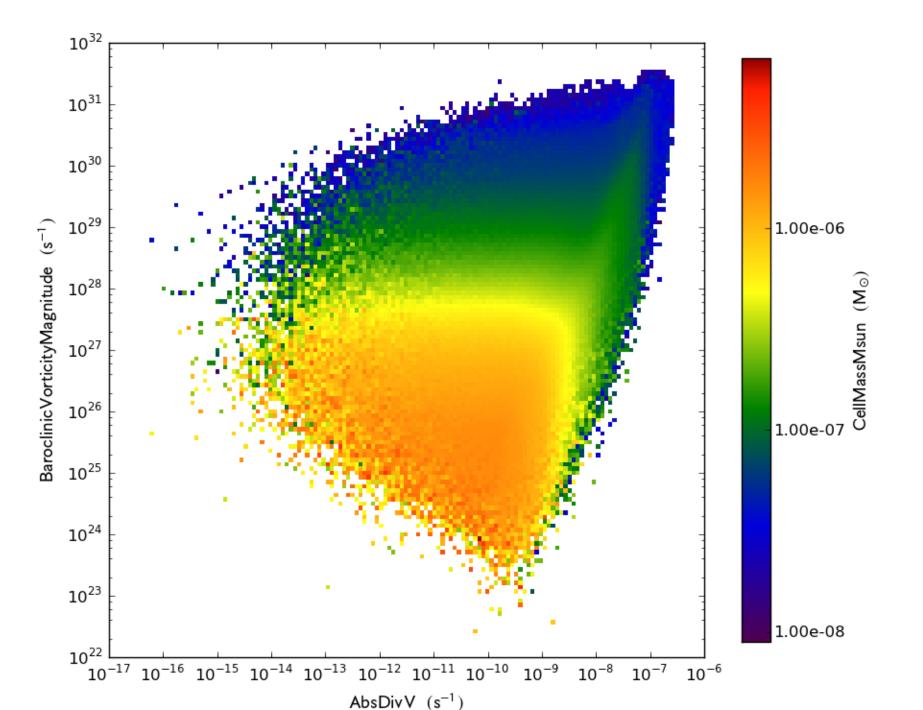
#### Increased inward velocity

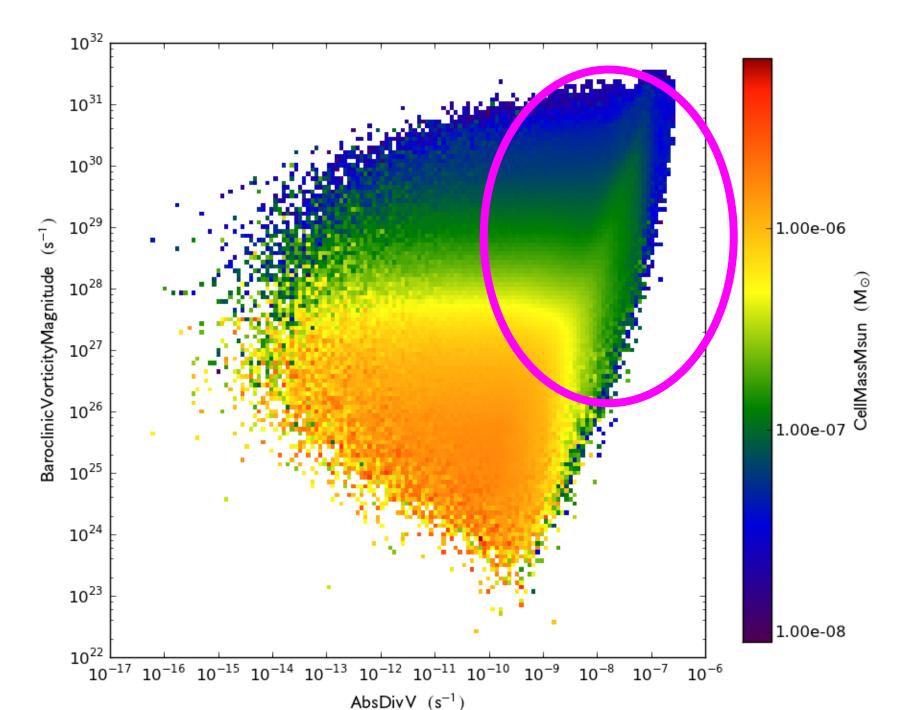
#### **Increased core temperature**

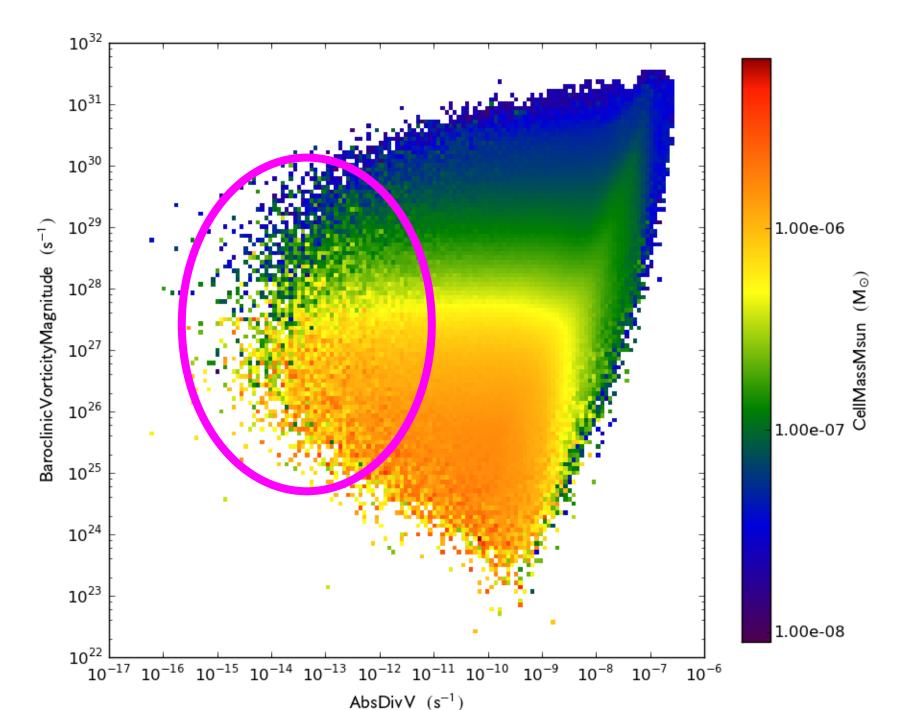
#### Decreased $H_2$ fraction



#### Shocks, chemothermal gradients, and convection.







# Disk structure is changed, but perhaps not by the magnetic field

# Dynamo action operates in cosmological initial conditions



#### **Collaborators**:

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